Swift Observations of Novae

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and the nova community at

http://www.swift.ac.uk/nova-cv/

Novae: SN1a progenitors?

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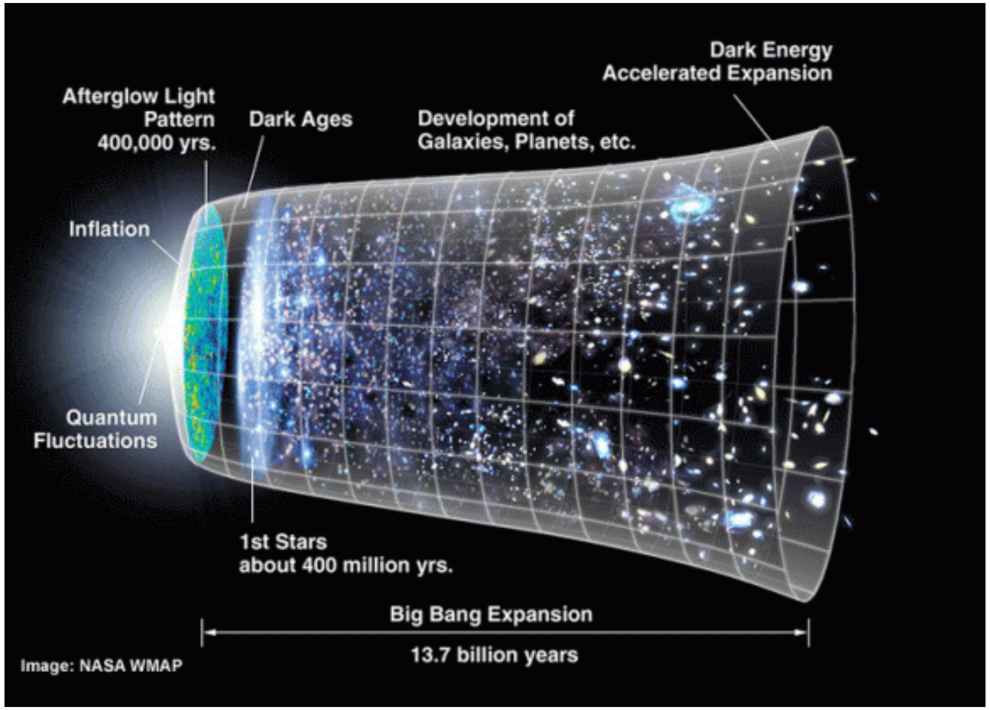
SN1a are key to measuring the acceleration of the Universe

- Thermonuclear explosion of a CO WD
- Progenitors are unknown (large literature!)
- Wang & Han 2012 review:
    - Single degenerate models:
        - WD+MS/RG/He
    - Double degenerate model:
        - WD+WD



- Gonzales-Hernandez+ 2012: lack of bright survivors means giant & subgiant companions excluded in SN1006
- Dilday+ 2012: PTF11kx circumstellar shells imply symbiotic nova prog.
- Broersen+ 2014: RCW86 is a remnant of a 1a that exploded in a 30 pc wind-blown cavity – requires an accretion wind, ie SD model
- Graur+ 2014: no HeII em before SN2011fe ➔ no high acc rate WD

Novae & X-rays

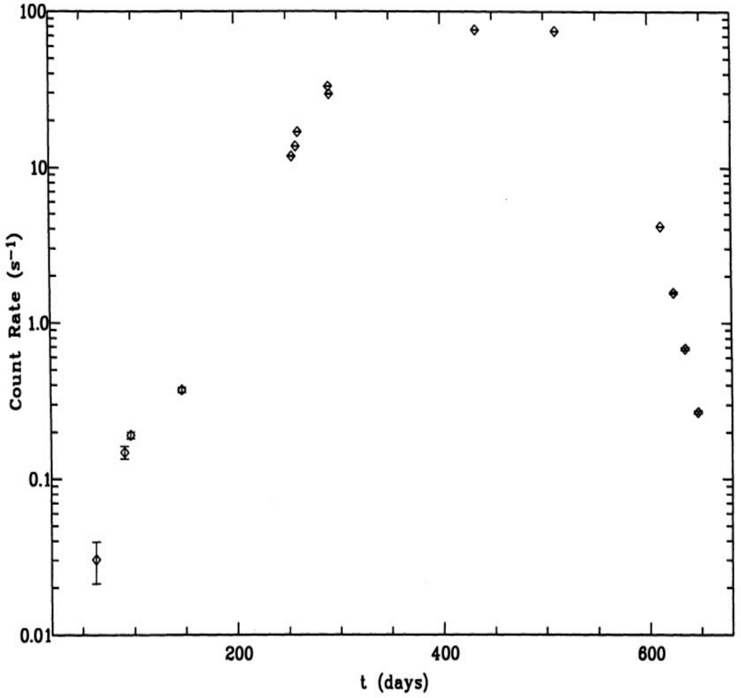
Novae:

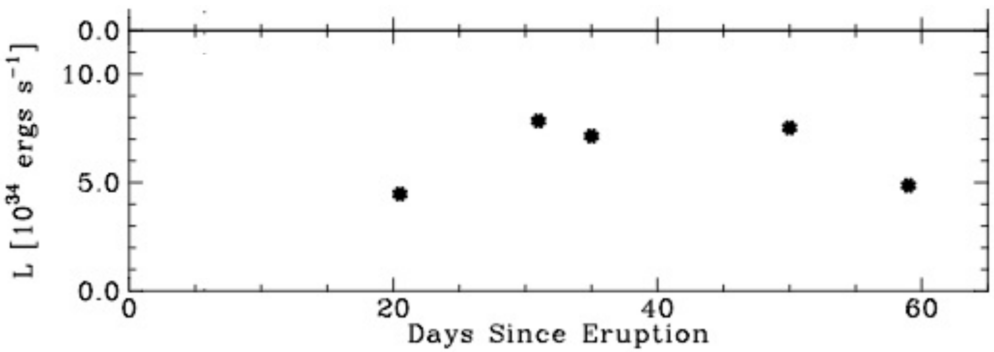
- Thermonuclear runaway burning of accreted material on a WD
- Orbital period can be hours to decades
- Recurrent novae have human timescale eruption cycle
- Unclear if M_{WD} grows or declines over many eruptions

Potential sources of X-ray emission from novae:

- Thermal emission from hot white dwarf
  - shock breakout
  - residual nuclear burning after ejecta dispersal
- High velocity shocks
  - internal shocks within the ejecta
  - shock of ejecta with shell from previous nova or planetary nebula
  - shock of later fast wind with earlier slower wind
- Re-established accretion

Pre-Swift novae X-rays





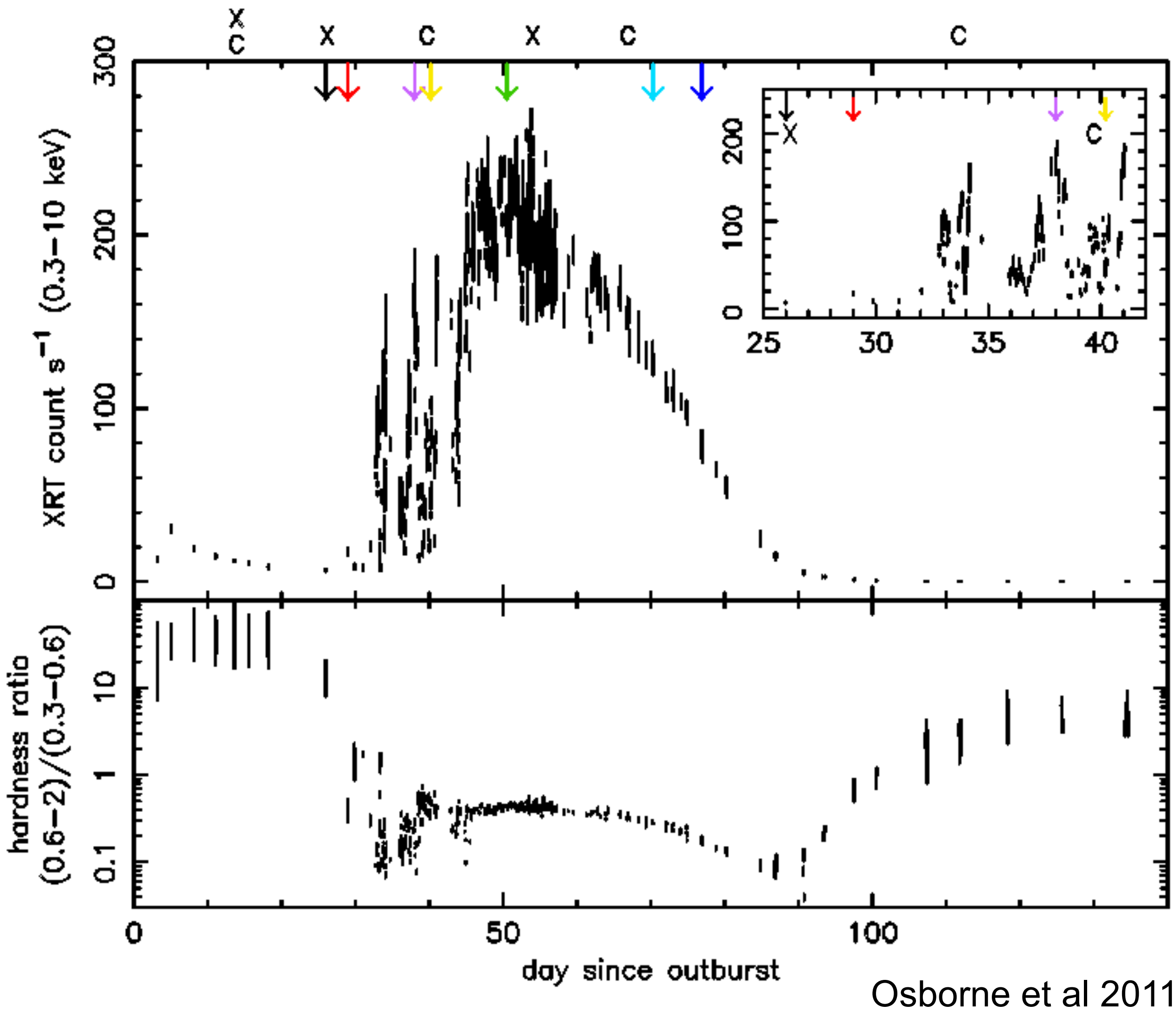
V1974 Cyg
16 ROSAT observations
Krautter et al. 1996

V382 Vel
5 ASCA observations
Mukai & Ishida '01

Swift novae stats

- Swift has observed 73 Galactic & MC novae within 11 years of outburst, of which:
- 43 detected in X-rays
- 12 novae have >100 ksec each: V1535 Sco, V745 Sco, V339 Del, N Mon 2013, T Pyx, U Sco, KT Eri, N LMC 2009, HV Cet, V2491 Cyg, V458 Vul & RS Oph
- Observations can start promptly: ~9 hrs from discovery for V339 Del (pre-nova for V2491 Cyg, U Sco, T Pyx & M31N 2008-12a)
- M31, M33 novae also observed

RS Oph

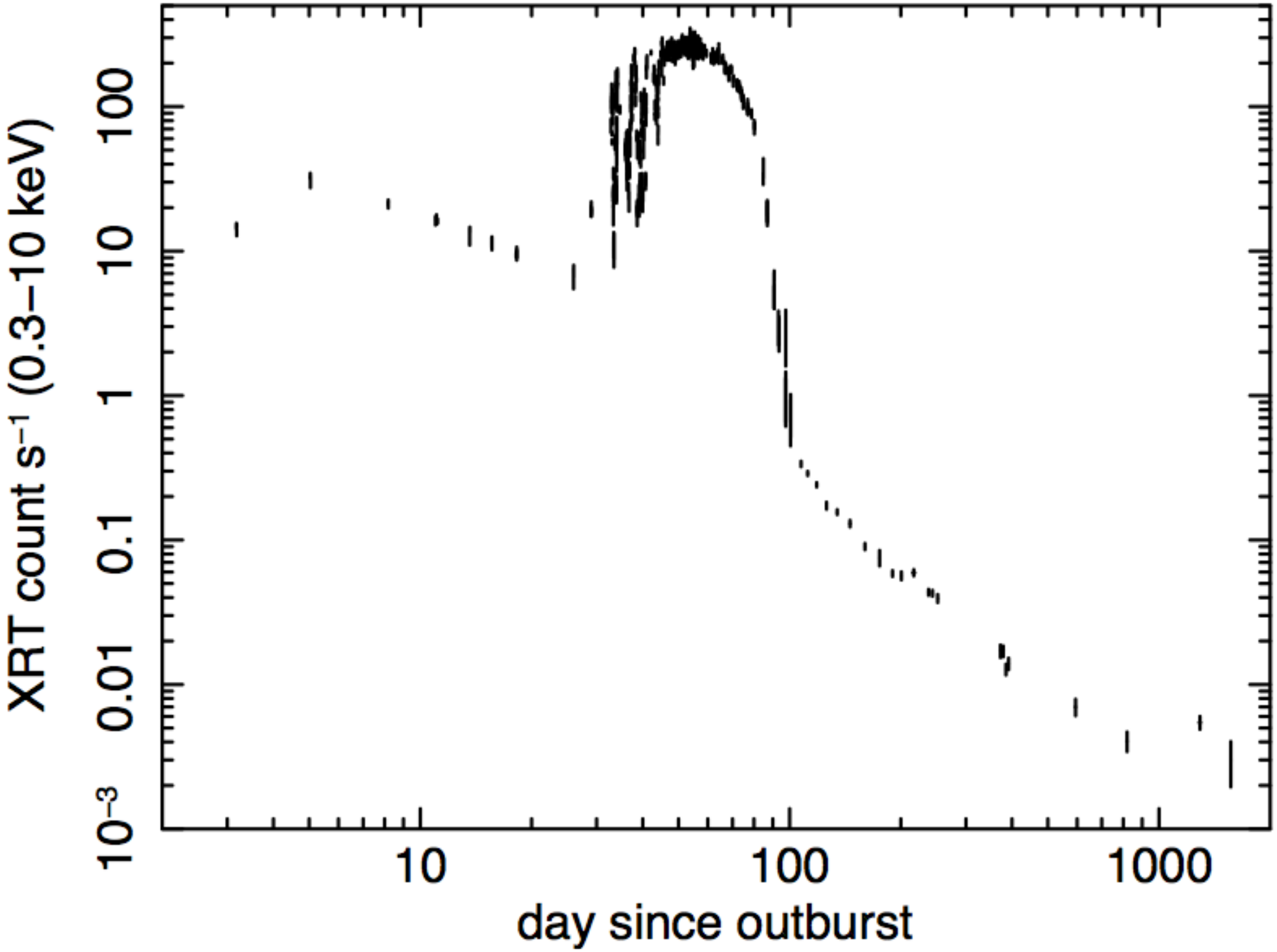


Osborne et al 2011

X-ray (0.3-10 keV) light curve shows:

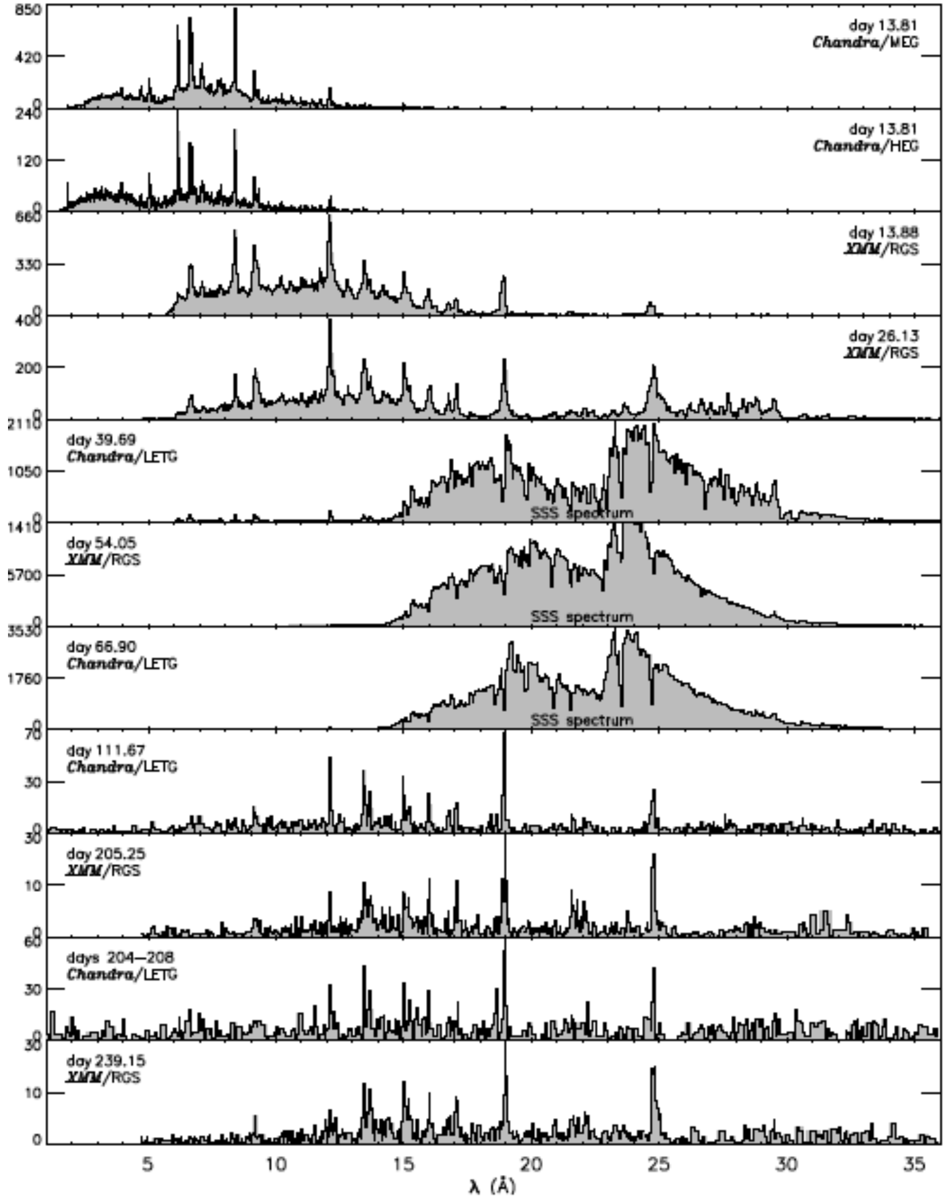
- Cooling hot gas emerging from red giant wind
- Noisy onset of super-soft phase, which lasts ~64 day in total
- Turnoff time → M_{WD}~1.35 $M_{☼}$

RS Oph



Osborne et al 2011

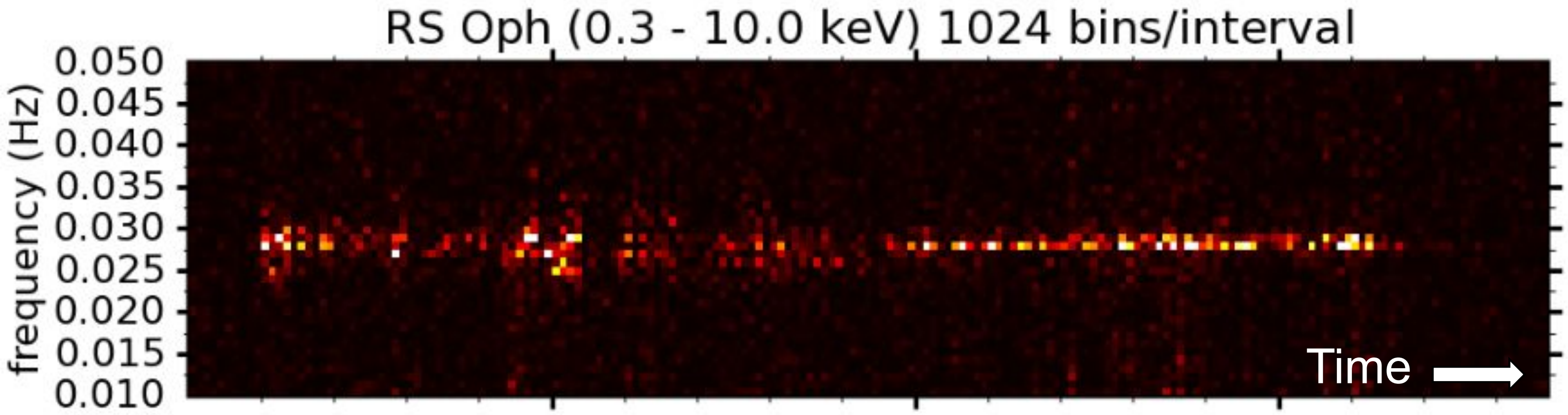
RS Oph

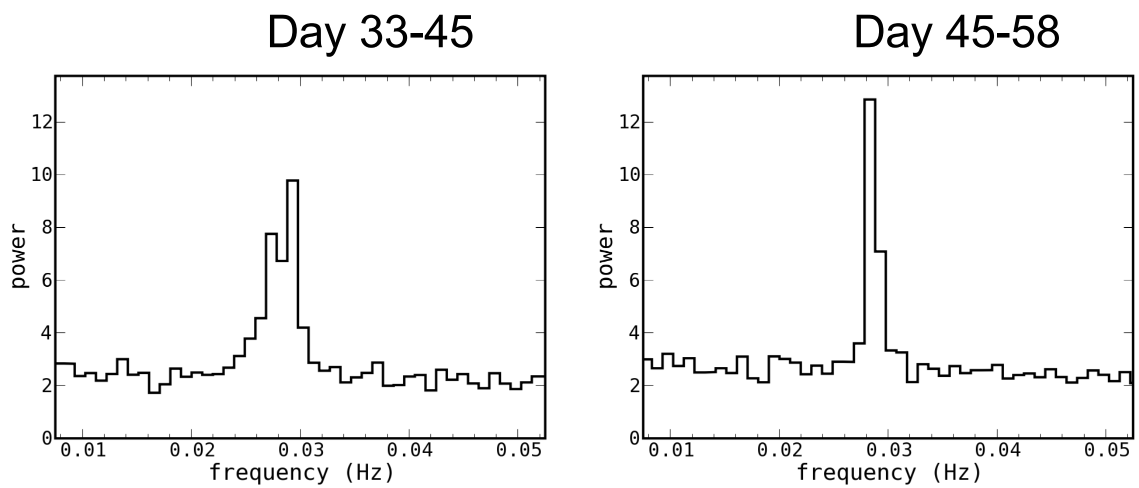


Ness et al 2009

RS Oph

A quasi-periodic modulation





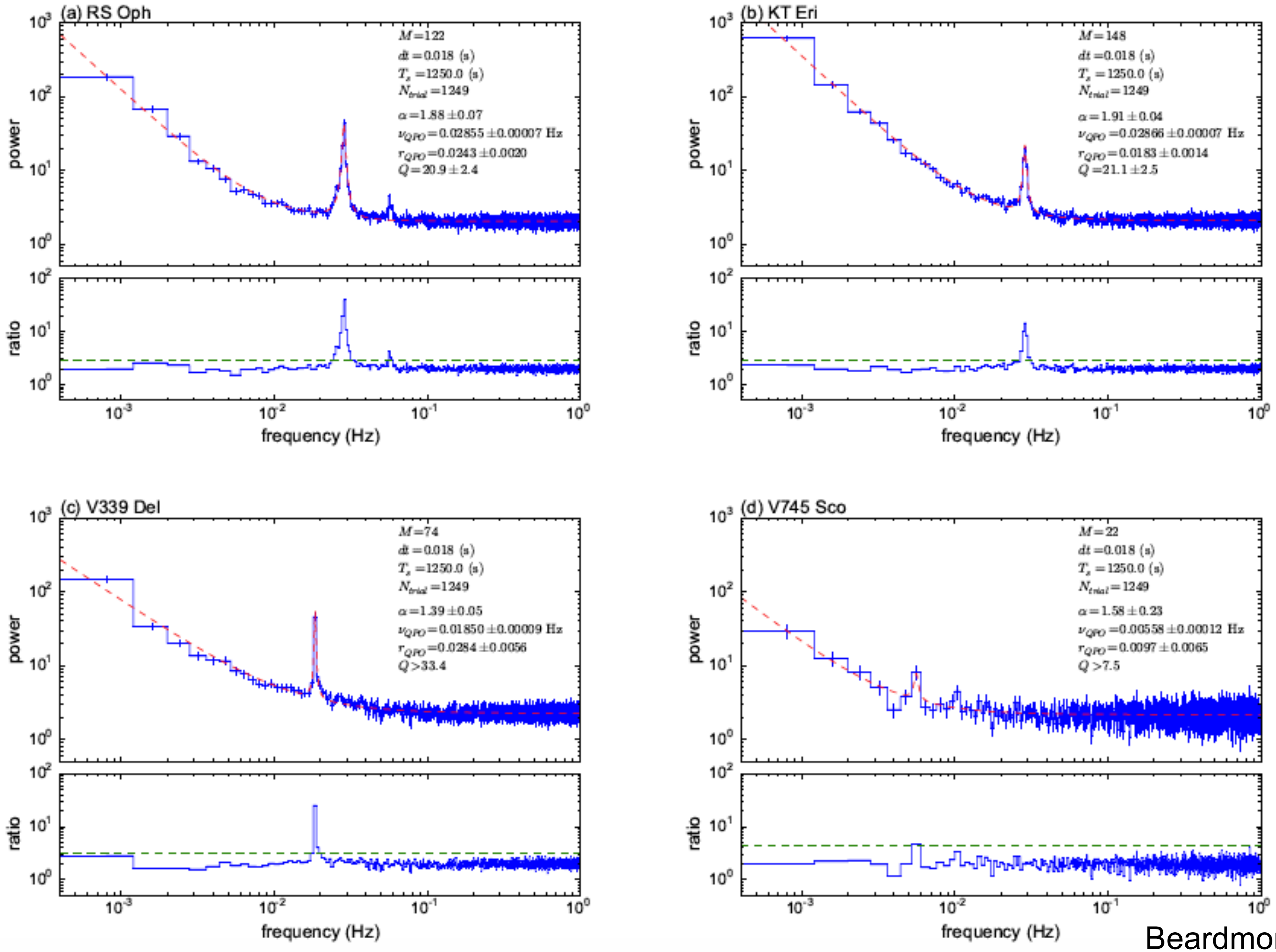
Osborne et al 2011

Period near 35s in soft X-rays between days 33-59

WD spin?

Nuclear burning instability?

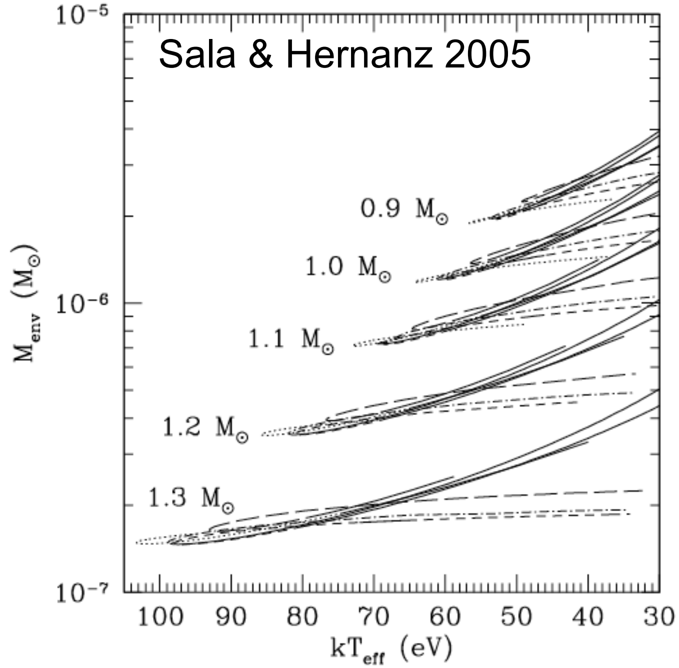
Super Soft novae QPO

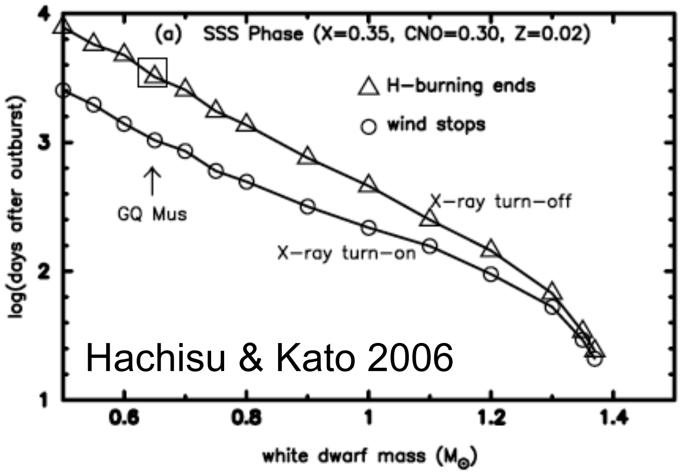


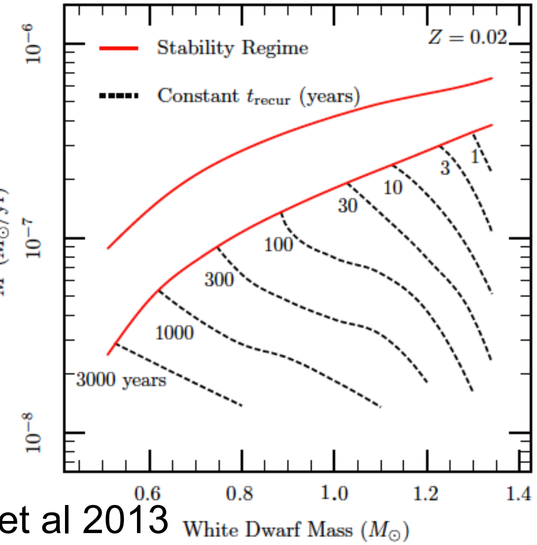
Beardmore et al 2015

What can we learn?

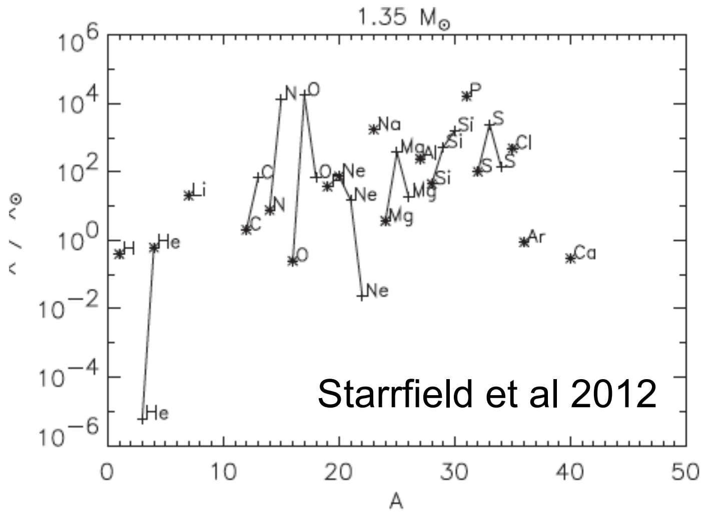
- kT_{max} gives M_{WD}
- T_{recurr} gives M_{WD}, acc rate
- T_{SS} gives M_{WD}
- Ejecta abundances give level of WD mixing, and so whether WD is gaining or loosing mass



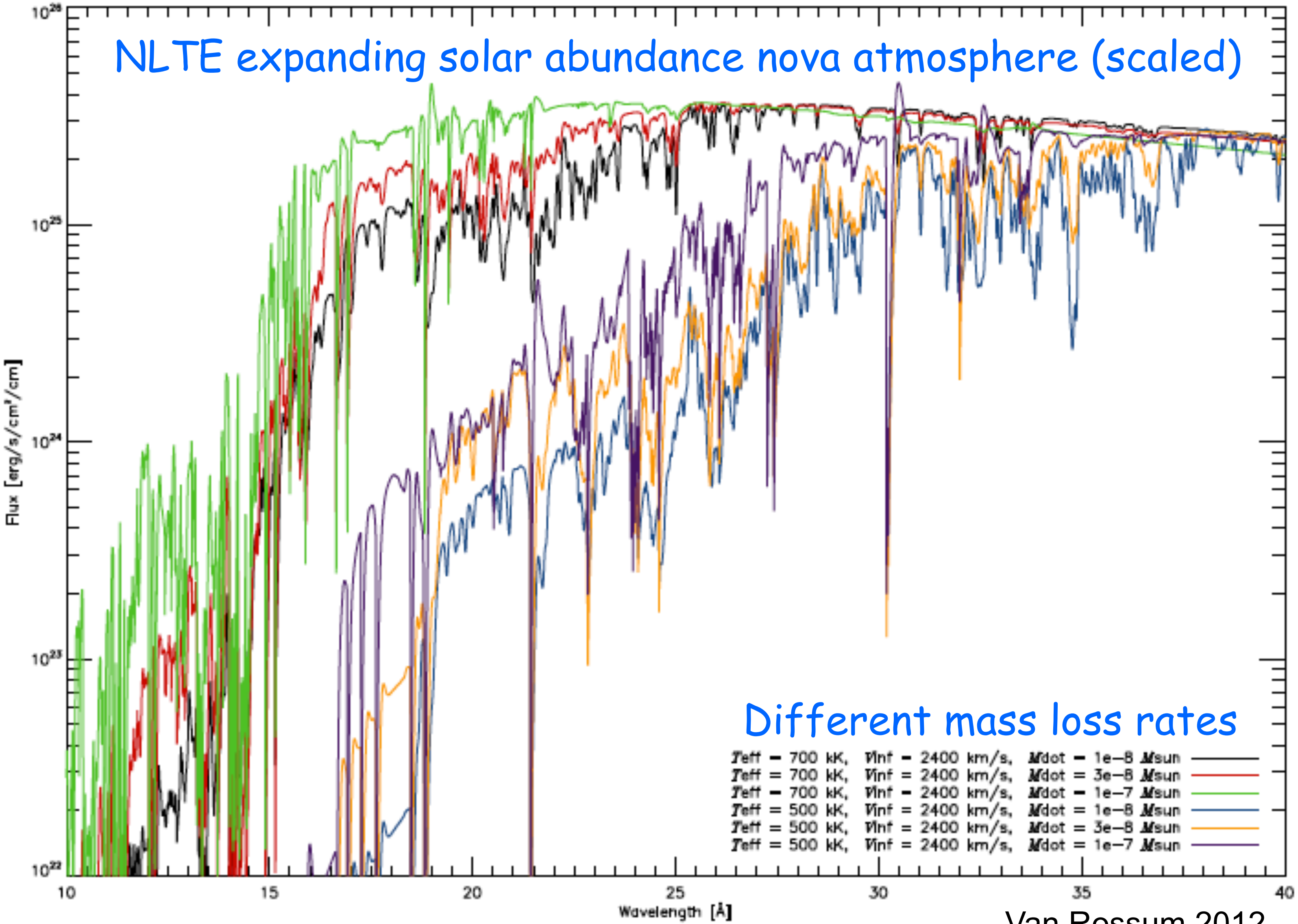




Wolf et al 2013

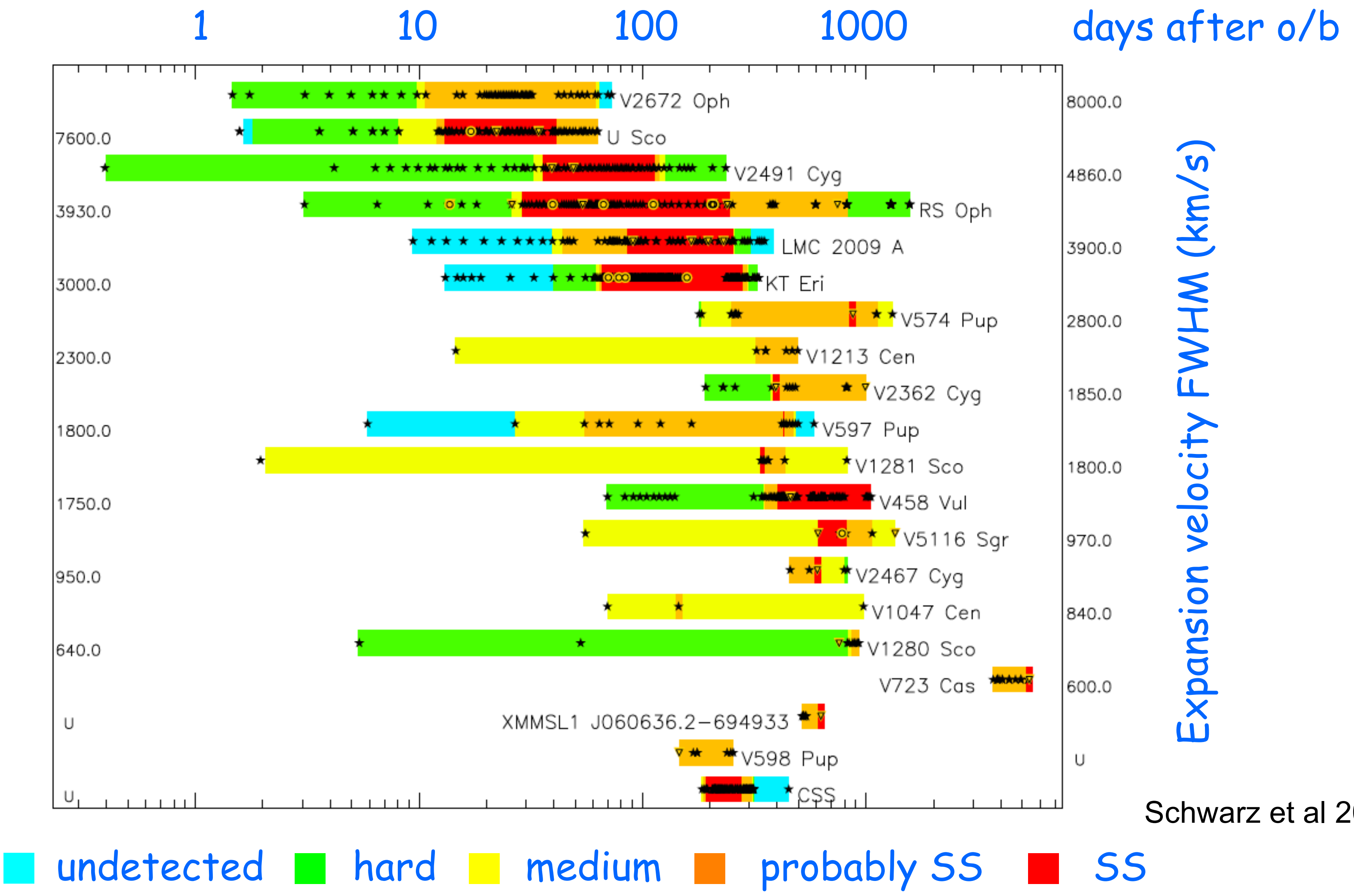


What can we learn?



Van Rossum 2012

Super Soft Swift Novae



Schwarz et al 2011

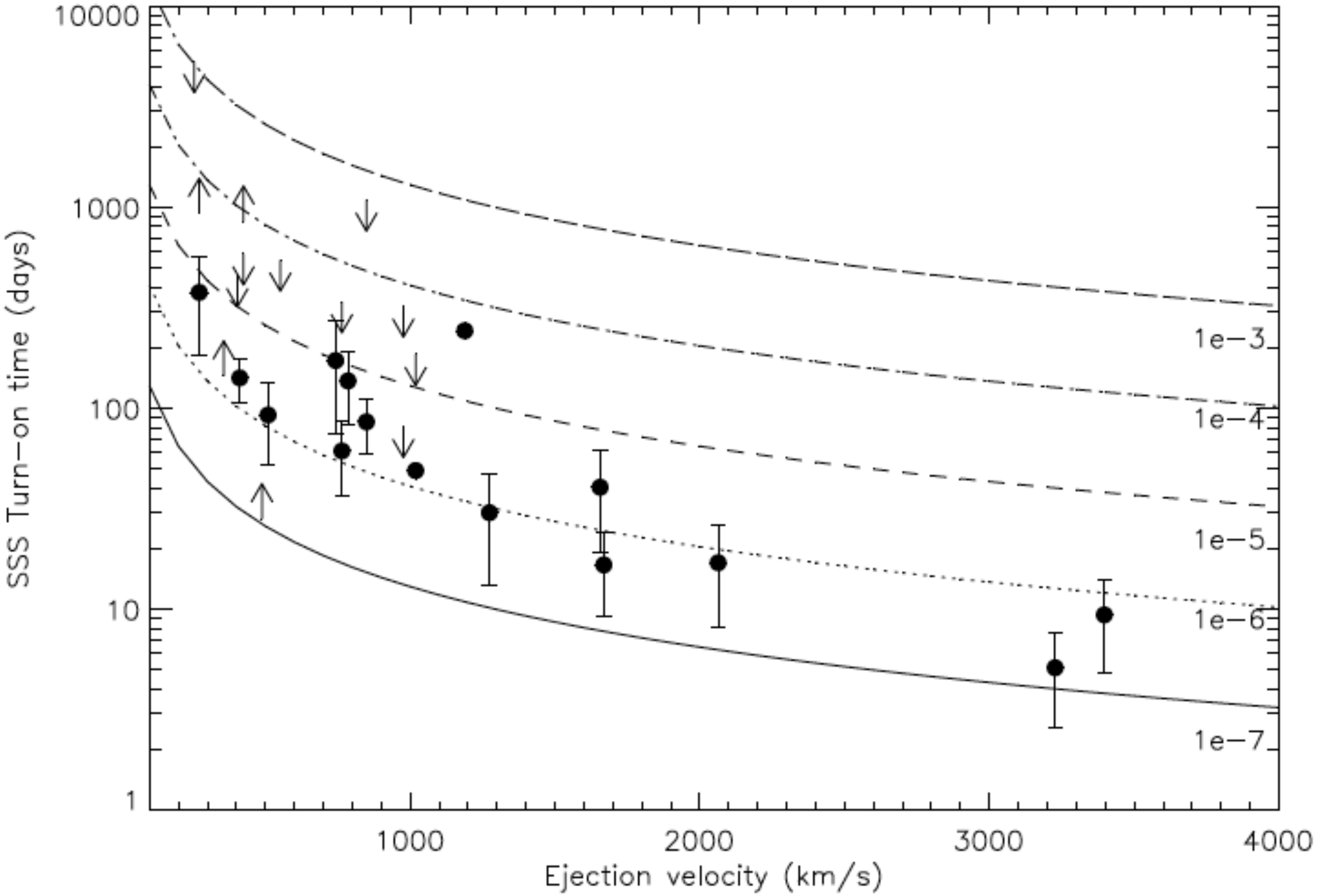
Super Soft Swift Novae

SS (& prob SS) defined as S = 2H

where S=0.3-1.0 keV c/s, H=1.0-10 keV c/s

- High expansion velocity ➔ high WD mass
- High expansion velocity ➔ early & short SS phase
  - Absorption & a strong hard component can be confusing
- The fastest novae have an early hard phase
  - Internal shocks in ejecta: $kT_{shock} = (3/16)\langle m \rangle V^2$
- Lack of SSS in previous samples due to observations being insufficiently early or late

SS novae turn-on times



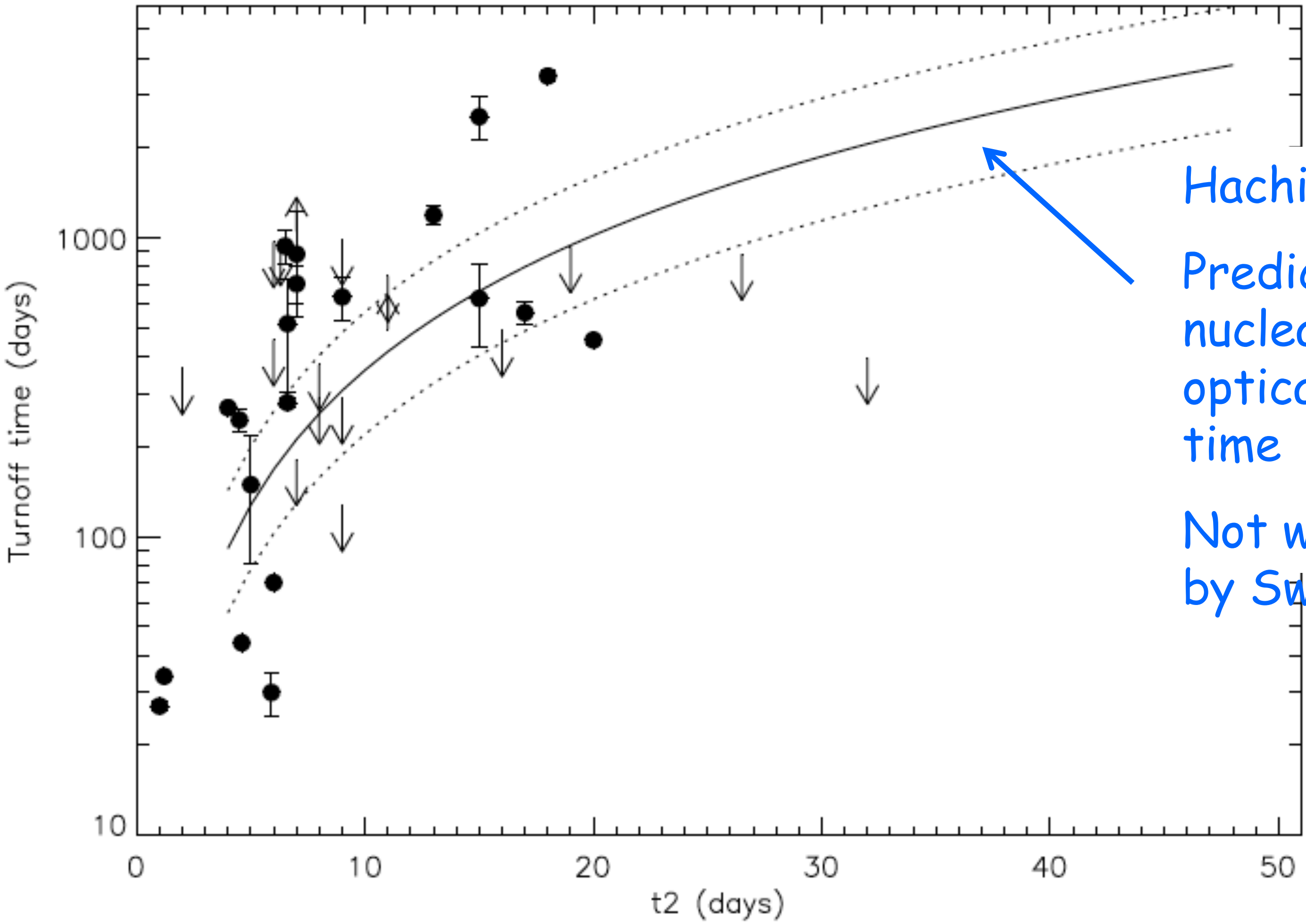
Ejected mass

The fastest novae eject the least mass

Schwarz et al 2011

They become SSS early as the ejecta thin out quicker

SS novae turn-off times



Hachisu & Kato 2010

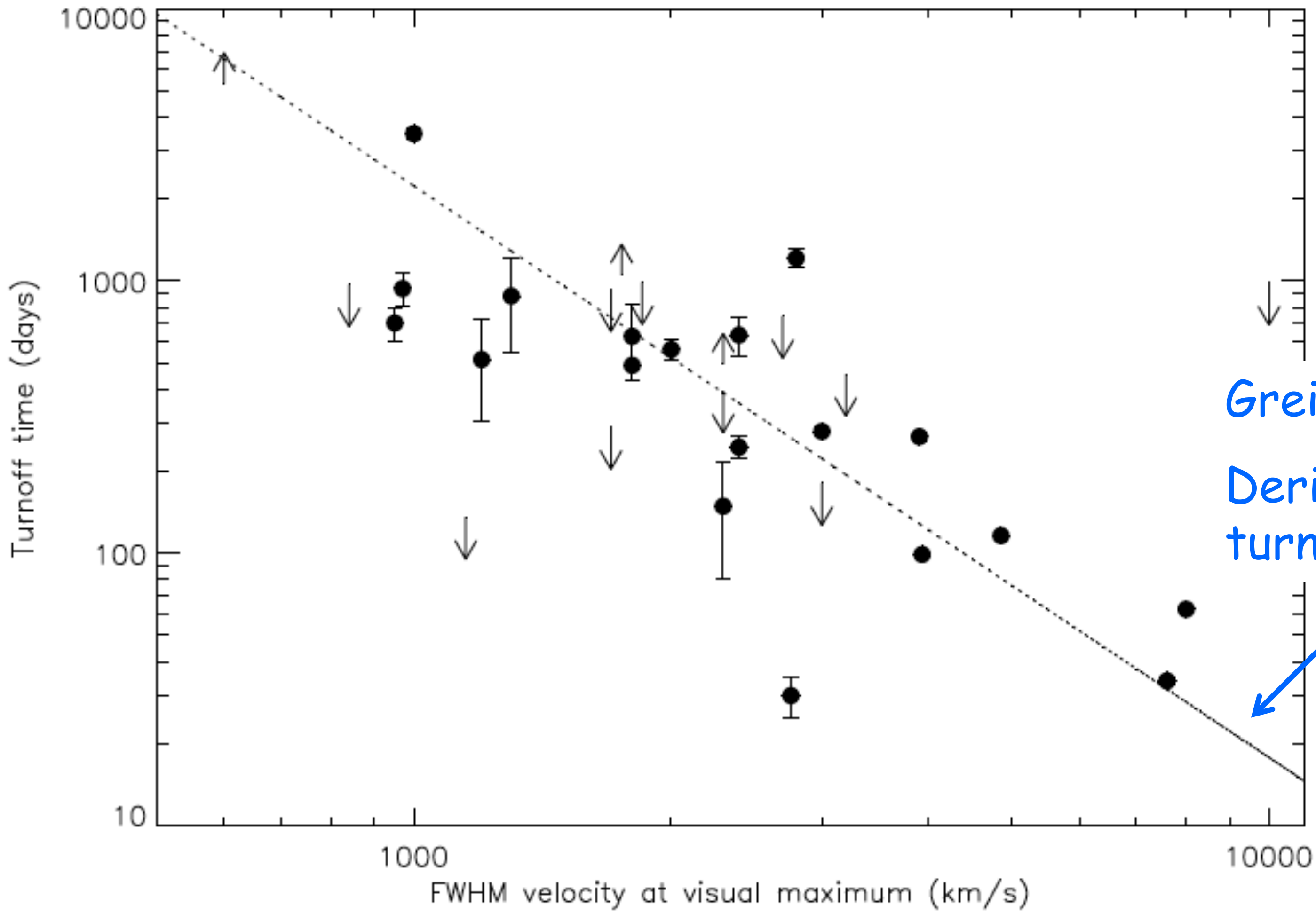
Predicted end of nuclear burning from optical LC break time

Not well supported by Swift data

Days to decline 2 magnitudes

Schwarz et al 2011

SS novae turn-off times



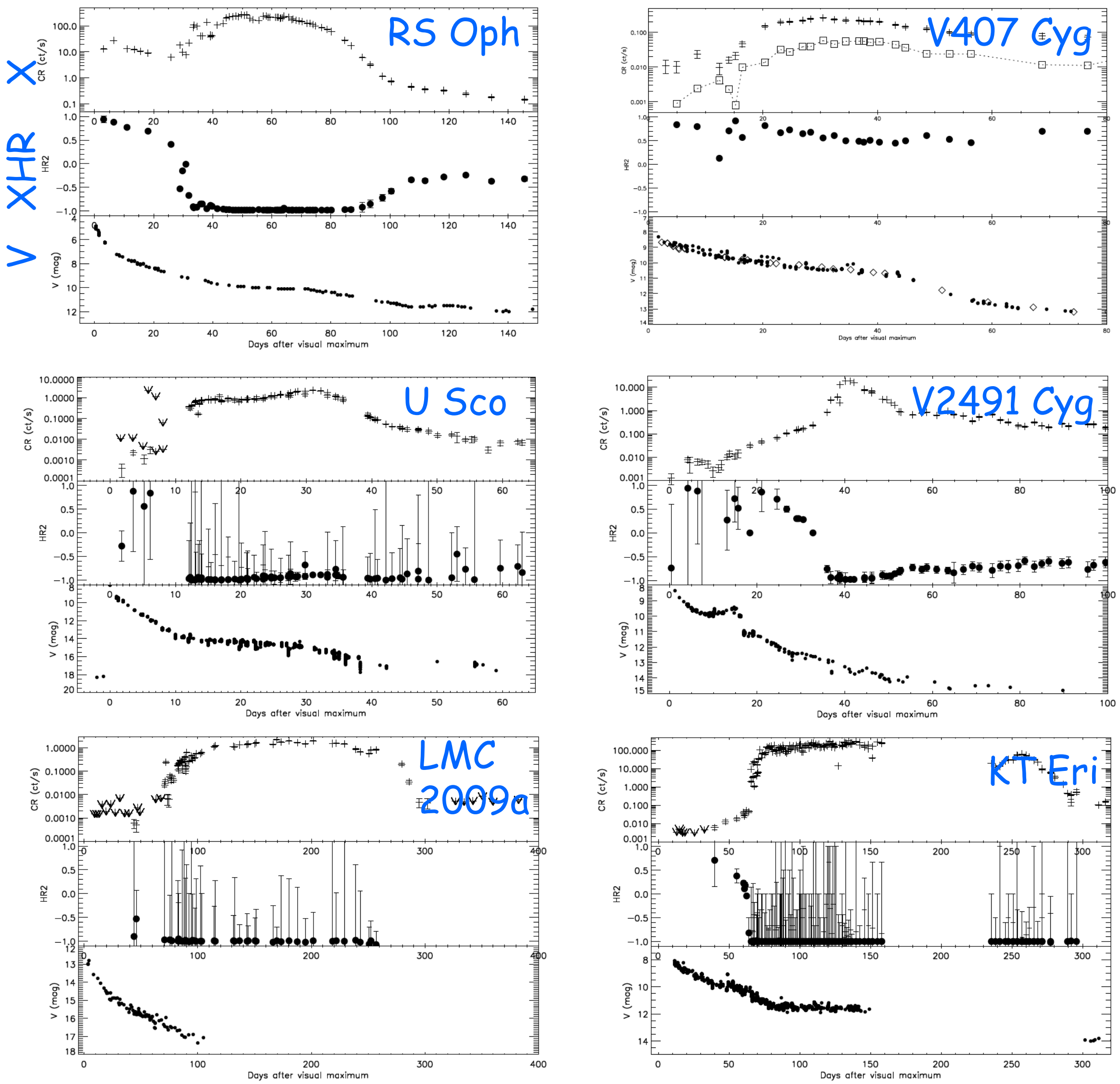
Greiner+ 03

Derived from 4 SS turn-offs

Ejecta velocity

Schwarz et al 2011

SSS = optical plateau?



Hachisu+ 08: "RNe optical plateau due to fading ejecta revealing irradiated accretion disk which ends when nuclear burning ends"

Our data agree with this, even for 2 of the 3 unconfirmed Rne

A proxy for SSS (as is [Fe X] 6375Å)

CSS 081007… = HV Ceti

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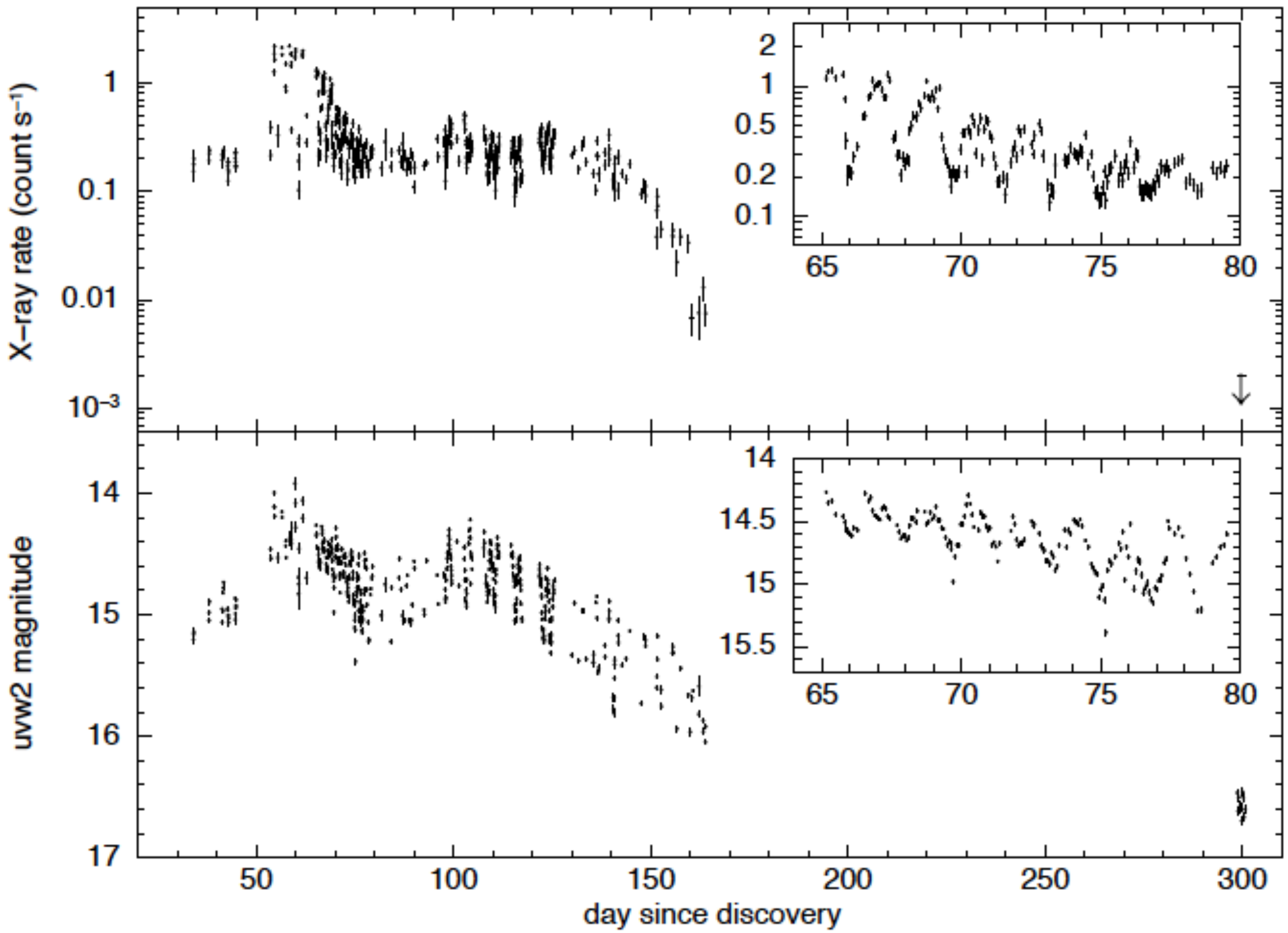
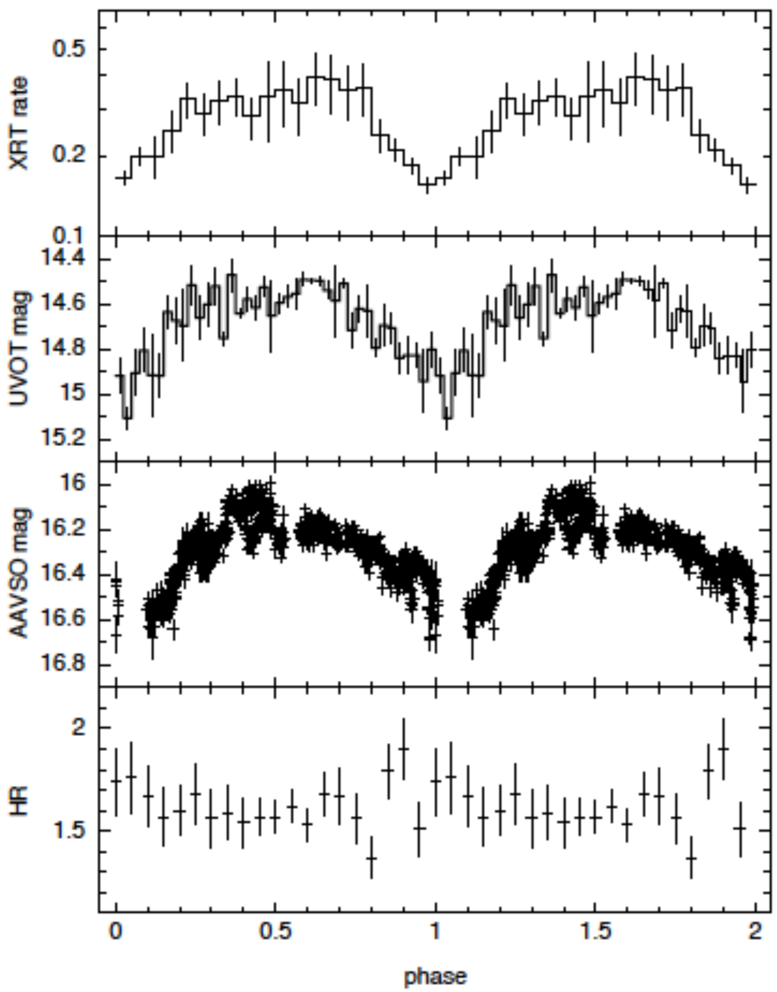
- Beardmore et al 2010: initial summary in AN
- Beardmore et al 2012: full paper - A&A 545, A116

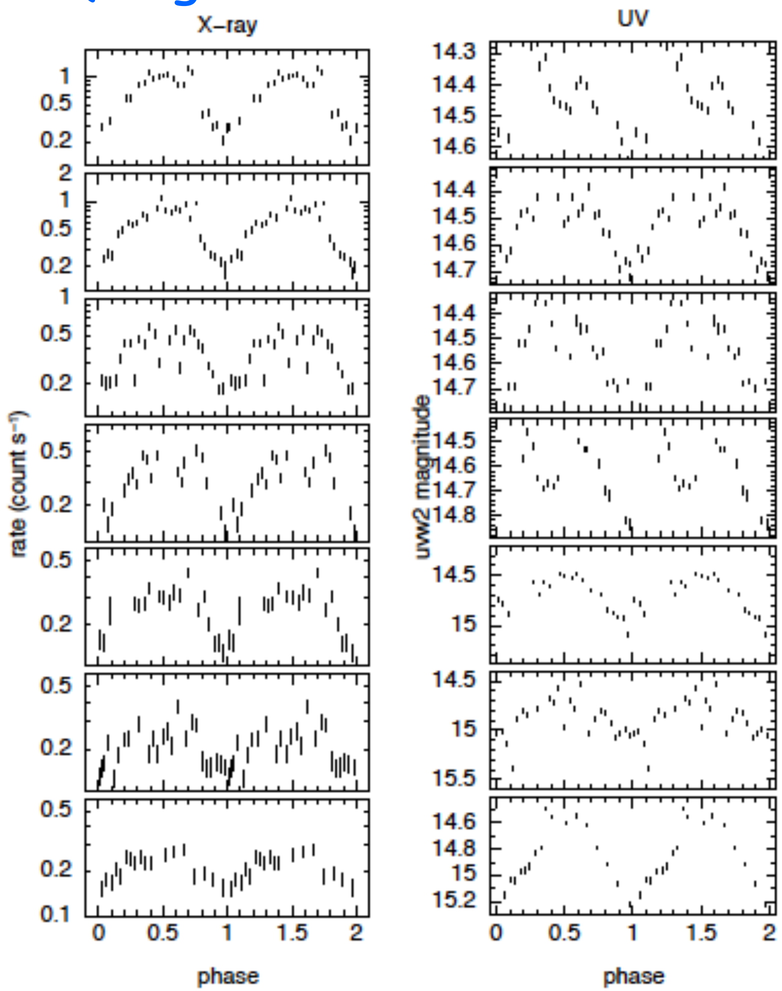


- ΔV > 4
- Ha vel ~ 1500 km/s
- Strong [Ne V]
- Pre-o/b rise ~ 1-2 mag
- Gal latitude = -44°
- Time of optical peak poorly constrained

HV Ceti

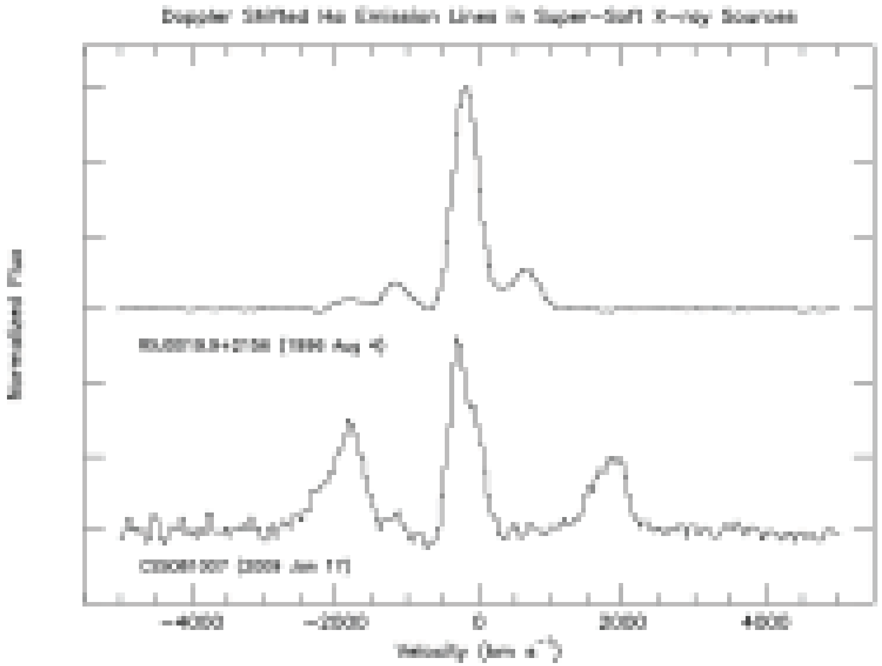
- 1.77 day modulation: orbital (cf GK Per) or poss precession
  - Broad modulation suggests large emission region
  - UV peak dips like SSS Cal 87 (bright inner accretion disk?)





HV Ceti

- Tri-peaked optical emission lines – which move
- Also seen in some Compact Binary SSS
- 'jet emission' / bipolar ejecta / accretion disk ??
- Hard to get both disk edge and high velocity jets in line of sight (unless they are broad)

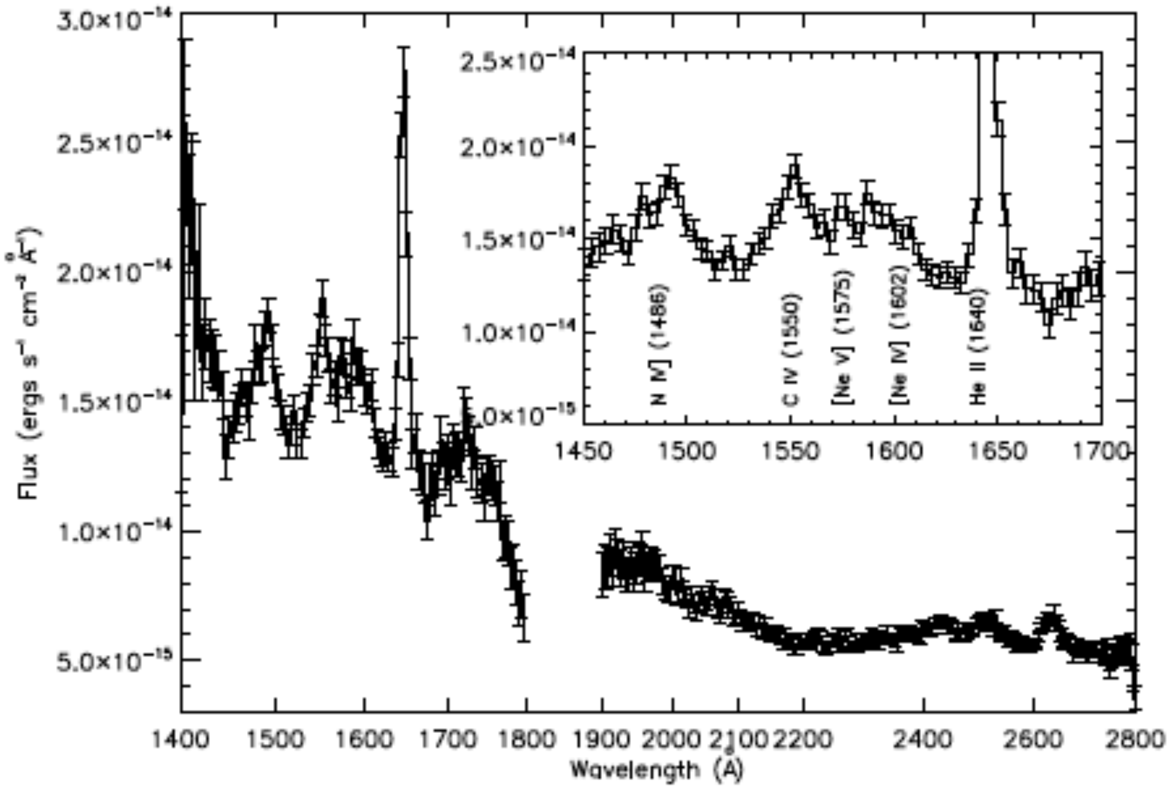


Wagner et al 2011

HV Ceti

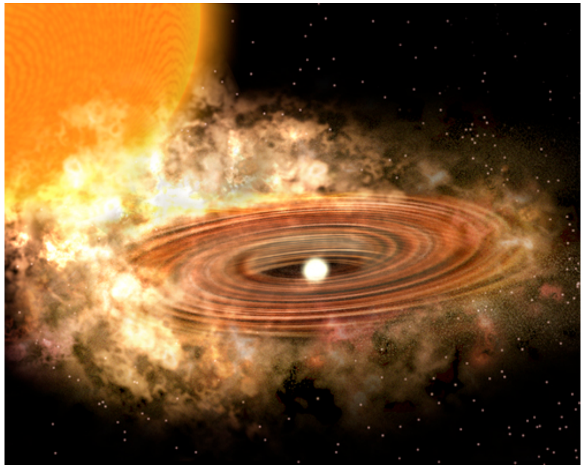
Galex spectra:

- no spectral variation
- few thousand times extrapolated X-ray spectrum
- but we know UV is modulated at 1.77 d like X-rays
- UV must come from inside accretion disk

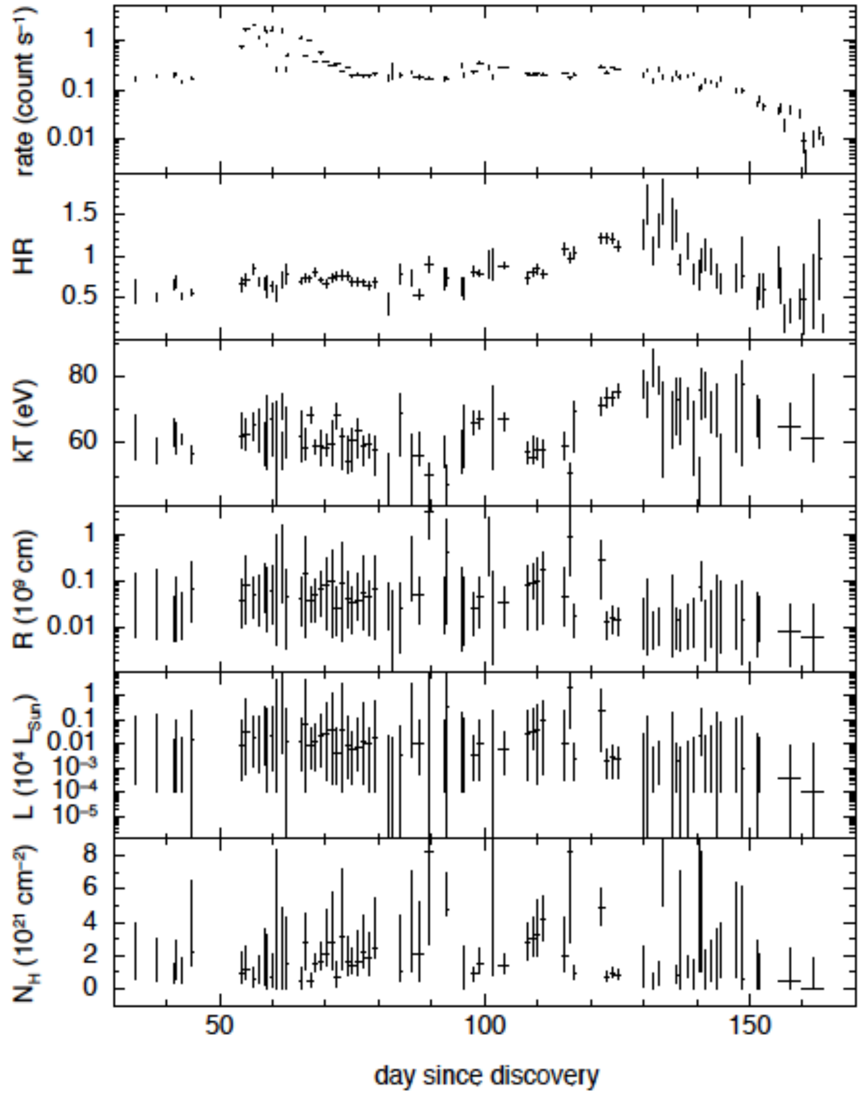


HV Ceti

•Suggests we see only scattered X-rays (τ<<1), while UV reflector sees hot WD directly



•Helps to explain R ~ few 10^7 cm from X-ray spectral fits (Rauch atmosphere model) – Well below expected ~10^9 cm



HV Ceti

- What about the trend in the periodic X & UV photometric variation?

    - X-ray max declines, min stays constant
while
    - UV max stays constant, min declines

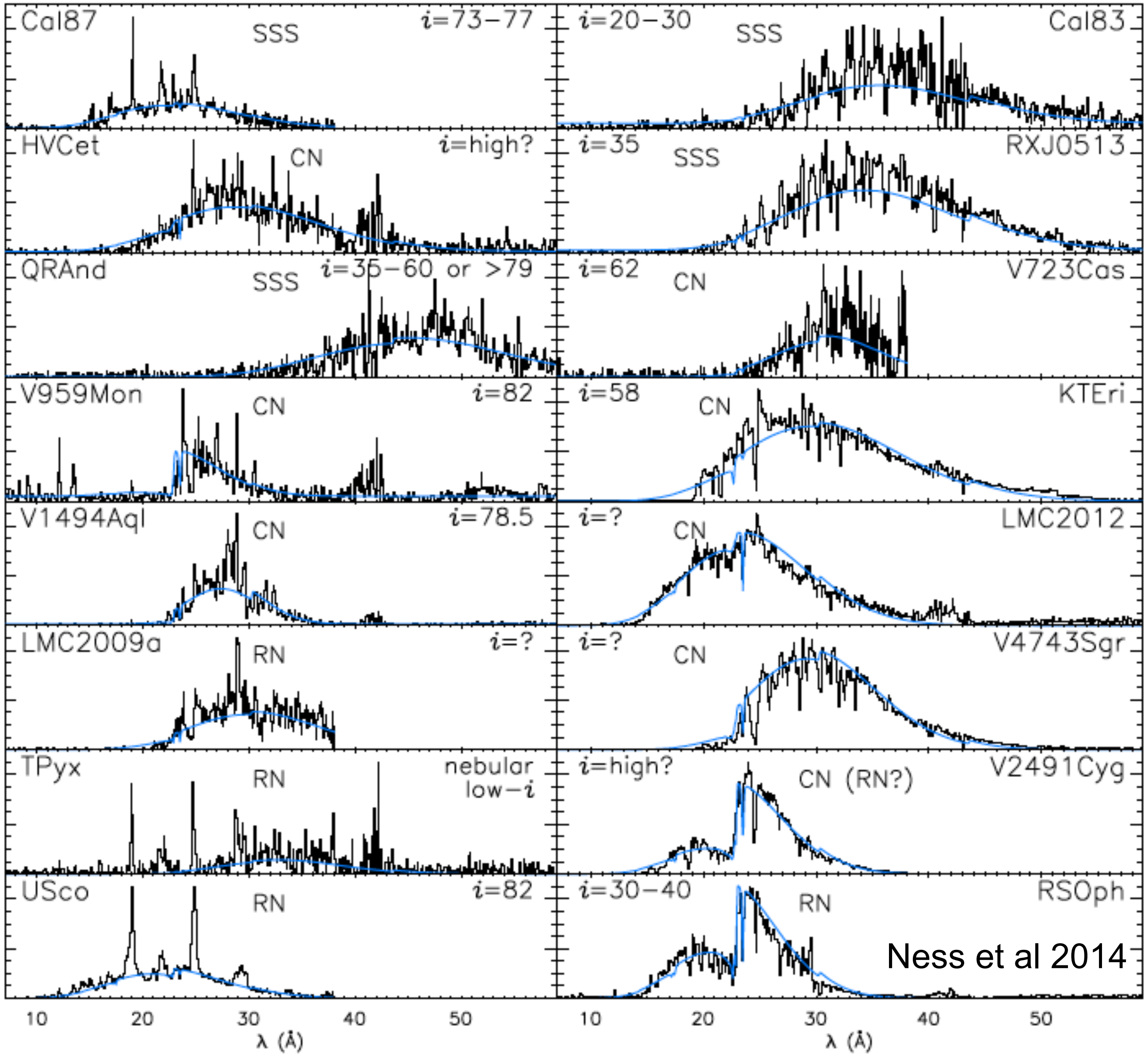
- Cannot be due to changes in disk rim height or size of inner scattering region

- No explanation to hand: worry about scattering cloud – UV reflection - disk obscuration concept?

SSS obscuration

Mostly novae

- High inclination SSSs are emission line dominated
- Low inclination systems dominated by continuum & absorption lines

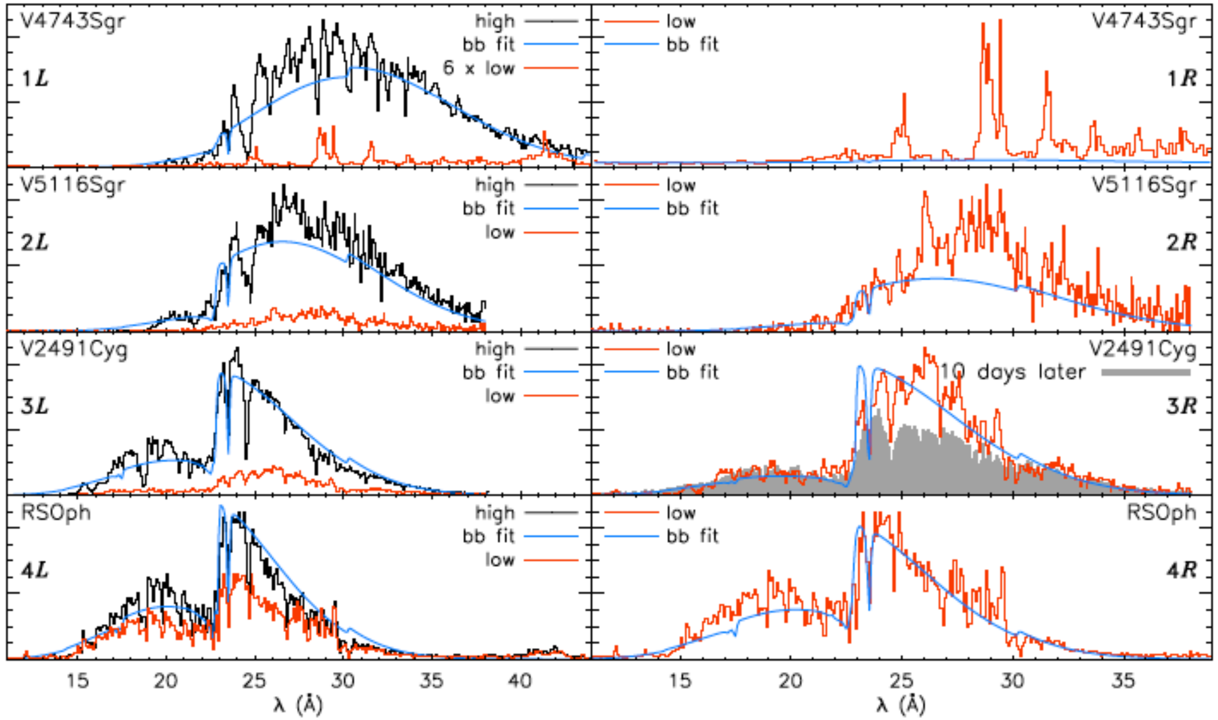


SSS obscuration

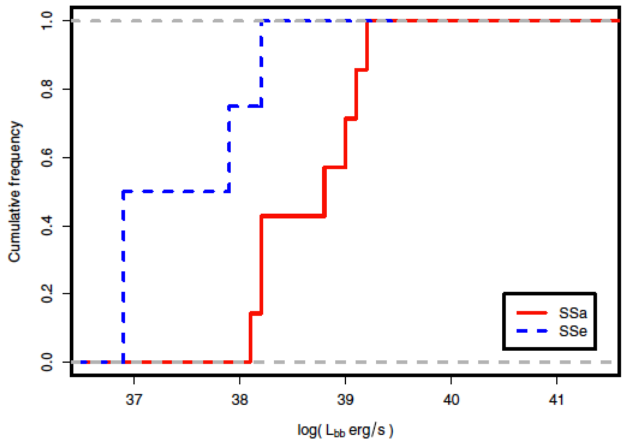
Sources are:

- continuum dominated when bright
- emission line dominated when faint

- More luminous sources are continuum dominated
- Less Luminous source are emission line dominated



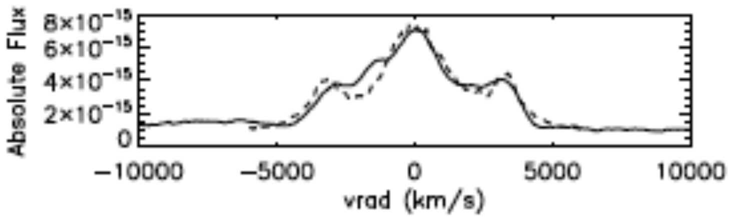
Ness et al 2014



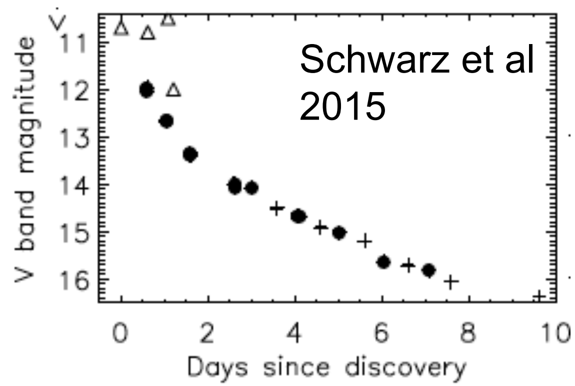
- Line of sight to WD is blocked at i>~70°
- Residual continuum seen via scattering
- Emission lines stronger where continuum is stronger ➔ photo-excitation
- Accretion disk exists at time of SS in novae

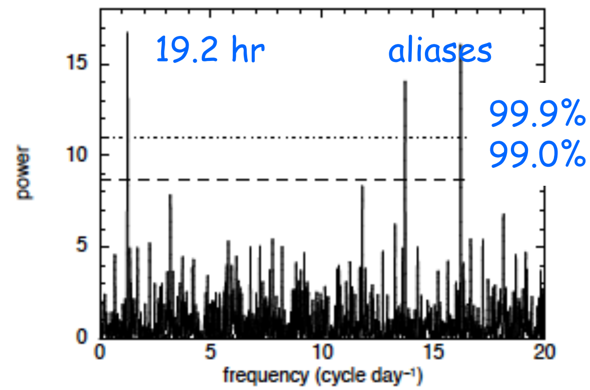
Nova LMC 2012

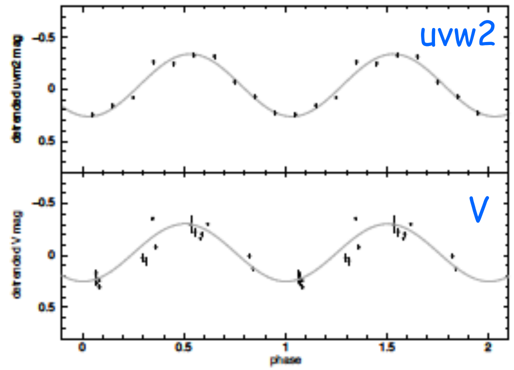
- Very fast decline: t_2 ~ 2 days (~ fastest ever seen)
- $P_{UV,\ opt,\ NIR}$ = 19.2 hr, not present in X-rays
- Emission line Chandra grating spectrum & optical emission line modelling suggests inclination ~55°



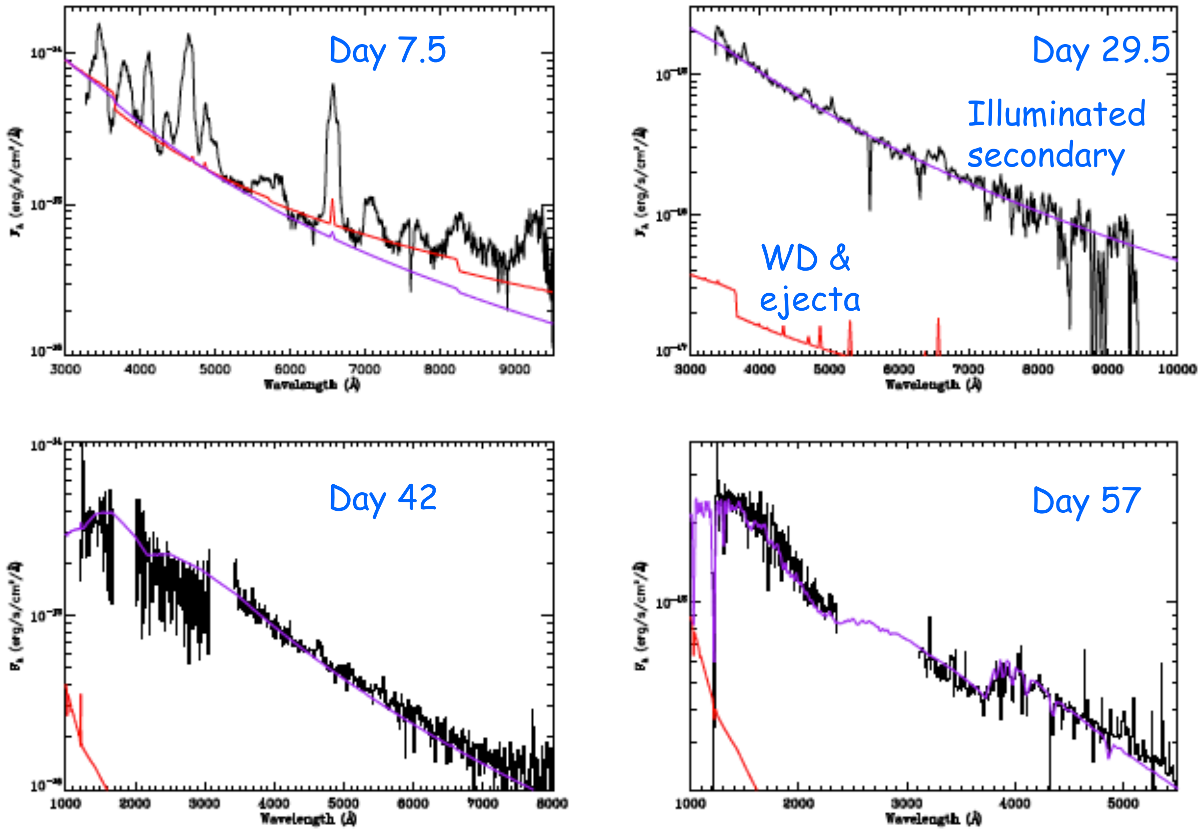
- Very low hard X-ray L (~1e31 erg/s) points to low M_{ej} (~1e-6 $M_\odot$)
- Short T_{SS} (= 50 d) at 1MK, high V_{ej} (~5000 km/s), low M_{ej} all point to a high M_{WD}







Nova LMC 2012

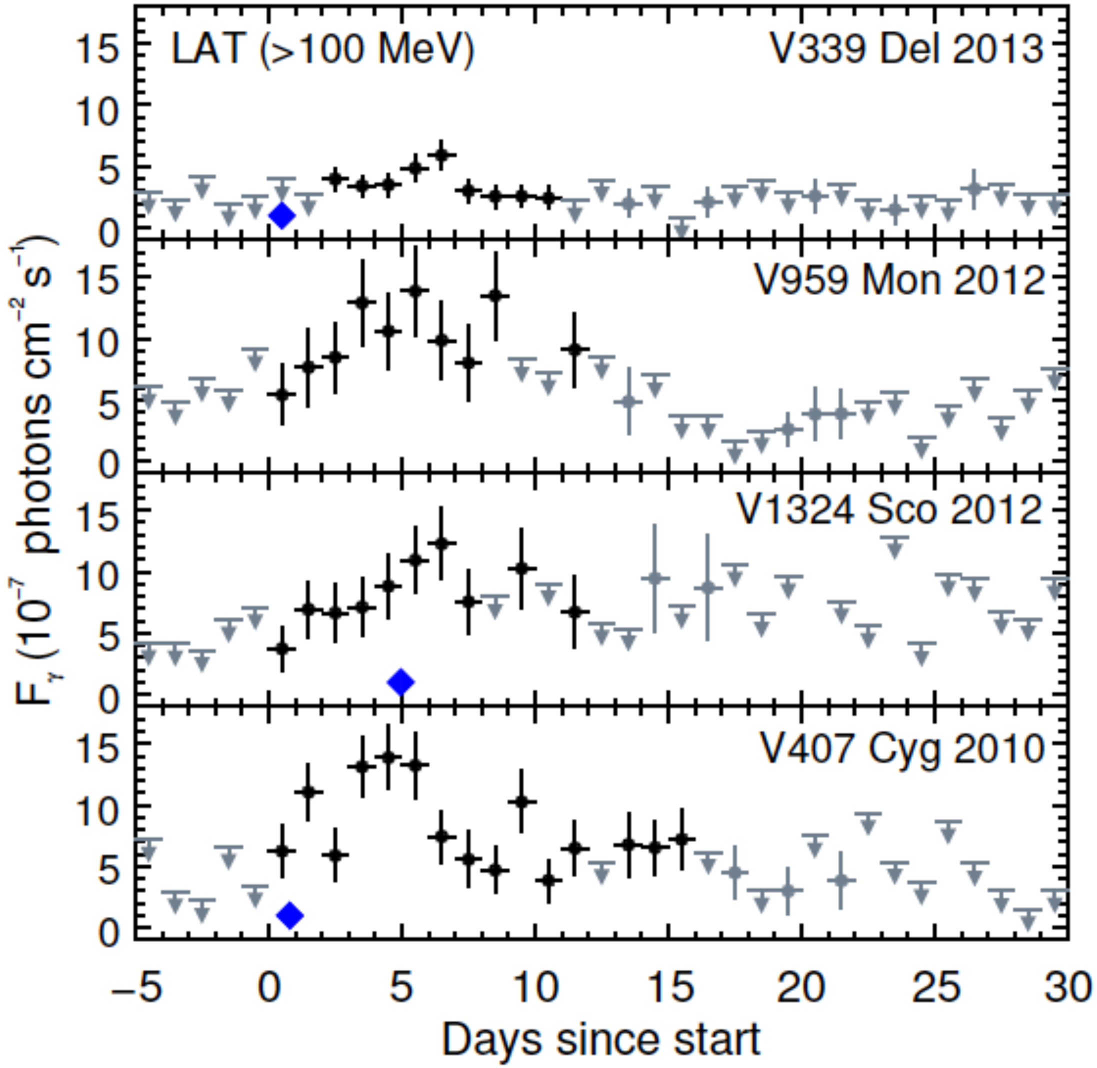


Schwarz et al 2015

- Optical/UV & X-ray spectroscopy constrain Rauch WD and Cloudy ejecta models, leading to the need for a 2nd UV/optical source:
  - Consistent with illuminated secondary, T ~ 20,000 K
  - Optical light curve *is not due to ejecta*
- Probable RN, gaining mass: SN1a progenitor – hard to find with t_2 ~ 2 days

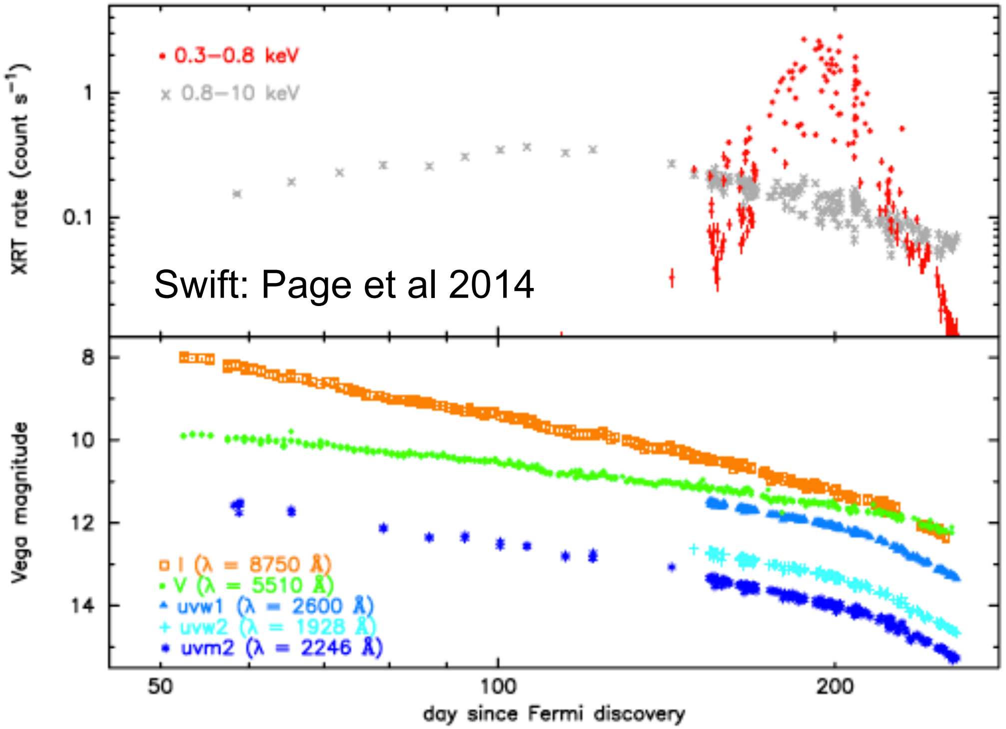
Fermi-LAT novae

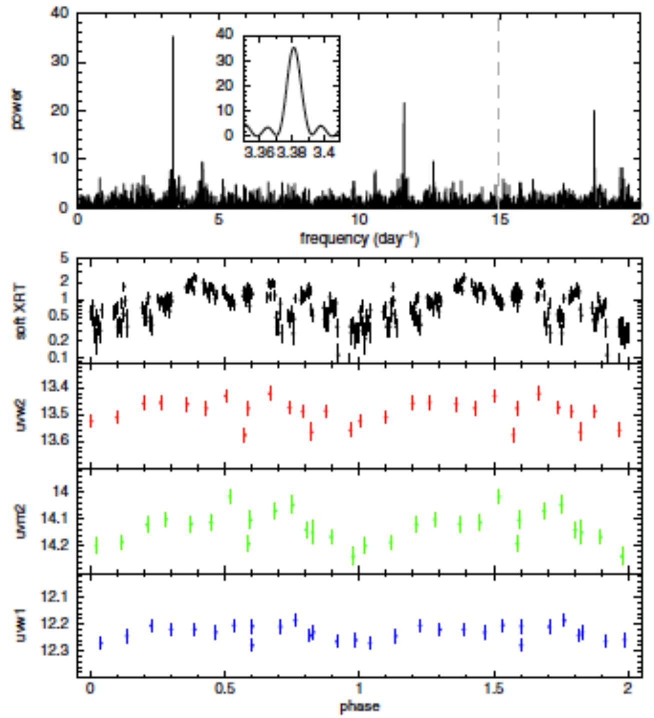
- 6 >100 MeV novae
- V407 Cyg: red giant (P_o~40-50y)
- N Sco = ?
- N Mon = K dwarf
- Emission peaks a few days after optical
- Discovery not widely expected: MeV line emission predicted, but not GeV continuum
- 1st had dense companion wind, good target for shock
    - photo-pion or IC origin
    - LAT novae would be rare
- No wind in N Mon 2012

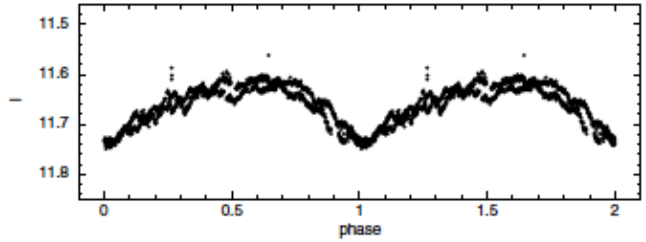


Cheung/LAT+ Science 2014
3σ detections marked bold. Optical peak blue

Fermi-LAT novae

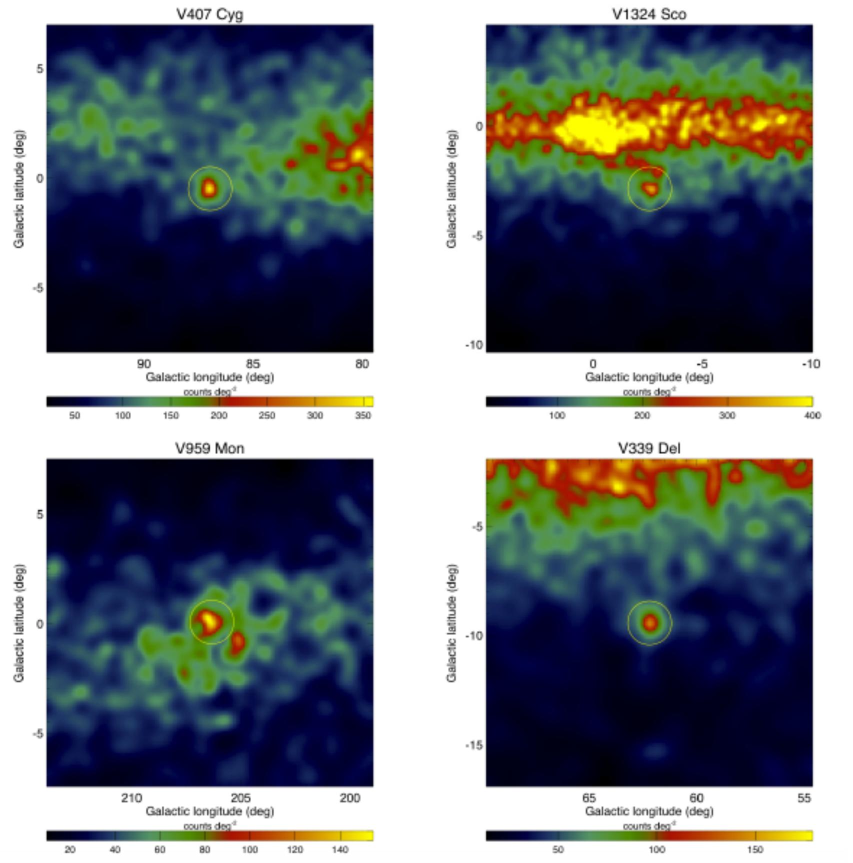






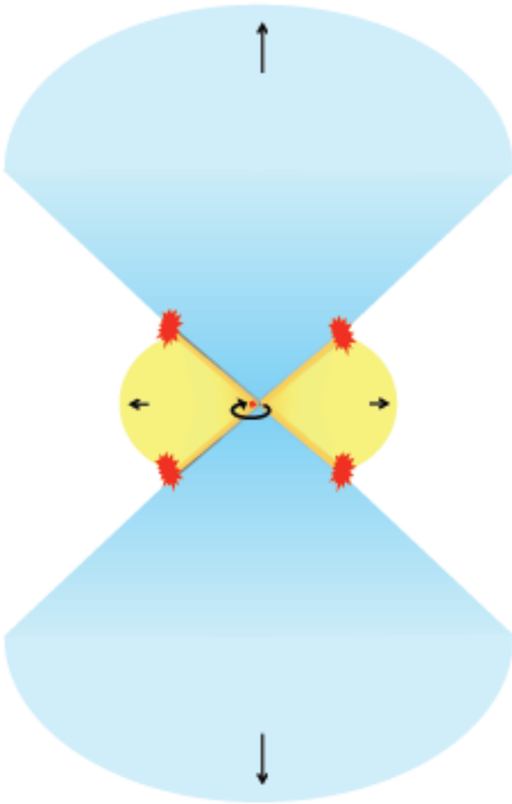
- N Mon 2012 seen in VHE-γ before optical
- 7.1 hr period in phase in X-ray, UV & I-band
  - Modulation due to disk rim obscuration?
- Presumed orbital period ➔ Msec ~ 0.8 $M_\odot$, rules out wind shock VHE-γ
- Martin & Dubus 2013 model: shock in gas around accretion disk

Fermi-LAT novae



Cheung/LAT+ Science 2014

V959 Mon

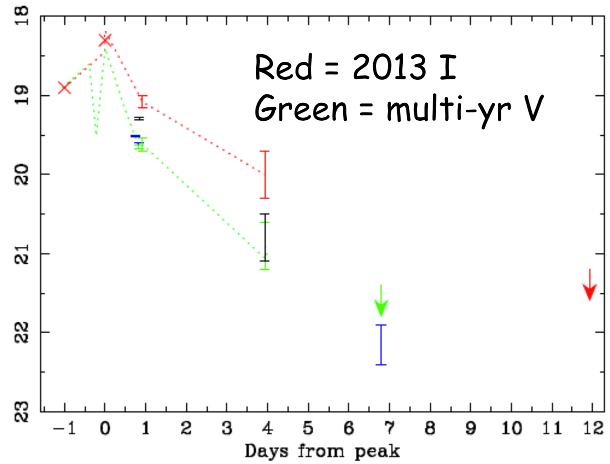


Chomiuk+ Nature 2014

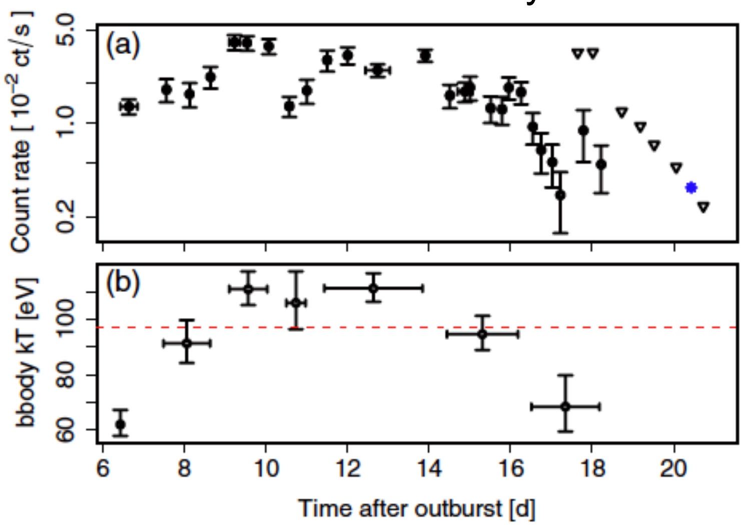
GeV emission as fast-late wind shocks on slower-denser material in orbital plane

M31 nova: T_{recurr} ~ 1 yr

- M31N 2008-12a outbursts in: 2008/12, 2011/10, 2012/10, 2013/11, 2014/10
- Shortest Galactic nova T_{recurr} ~ 8 yr (U Sco)
- Very fast V-band decline (t_2 ~ 4 days)
- Very high M_{WD} (~1.4 $M_\odot$) and accretion rate (log Mdot ~ -7 to -8)
- Very short SS phase seen by Swift (T_{SS} ~ 20 days) ➔ high M_{WD}
- T_{BB} ~ 97 eV ➔ very high M_{WD}
- Also found by Rosat as SSS in 1992/2 & 1993/1 and by Chandra HRC in 2001/9
- If this is a CO WD, it is a good SN1a progenitor candidate

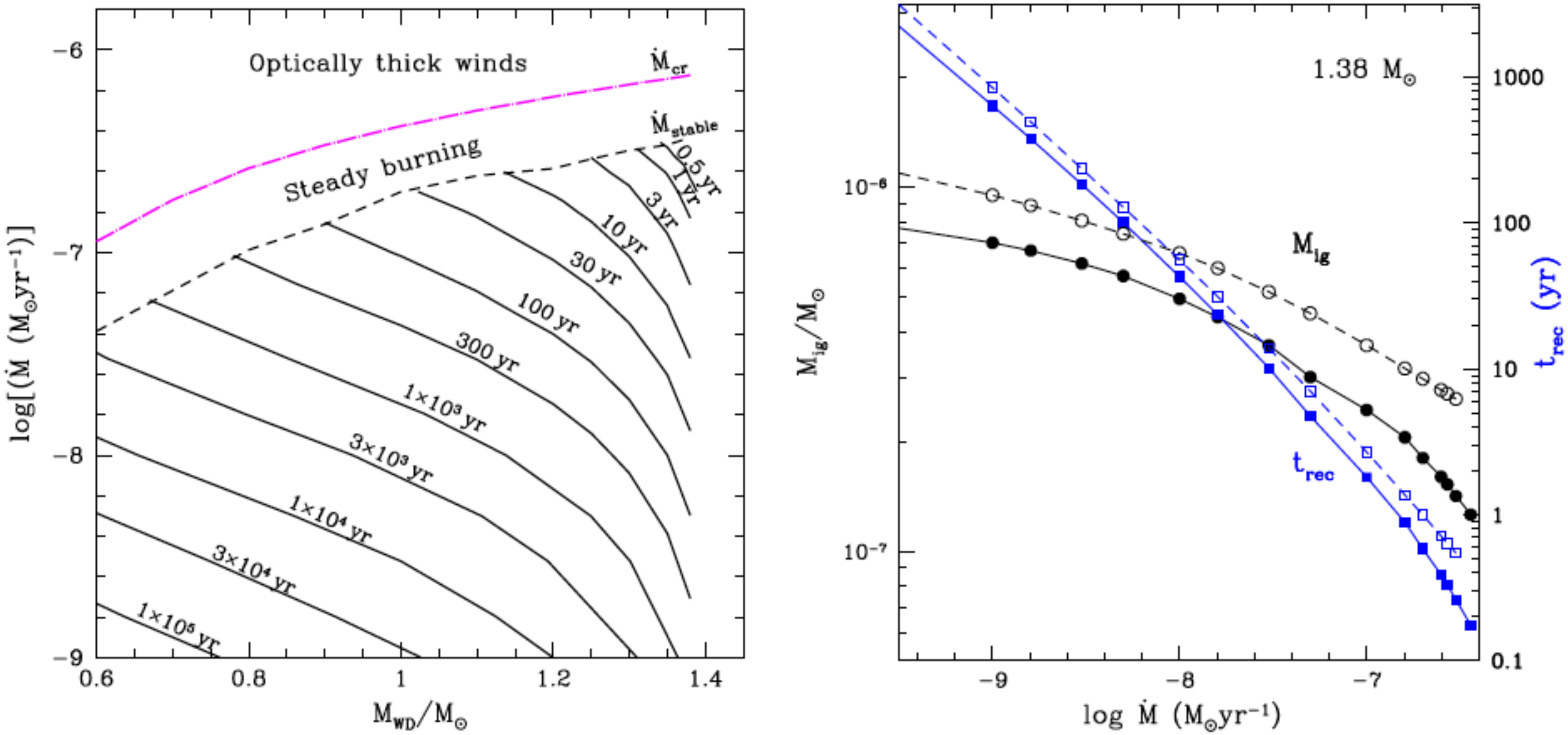


Darnley et al 2014



Henze et al 2014

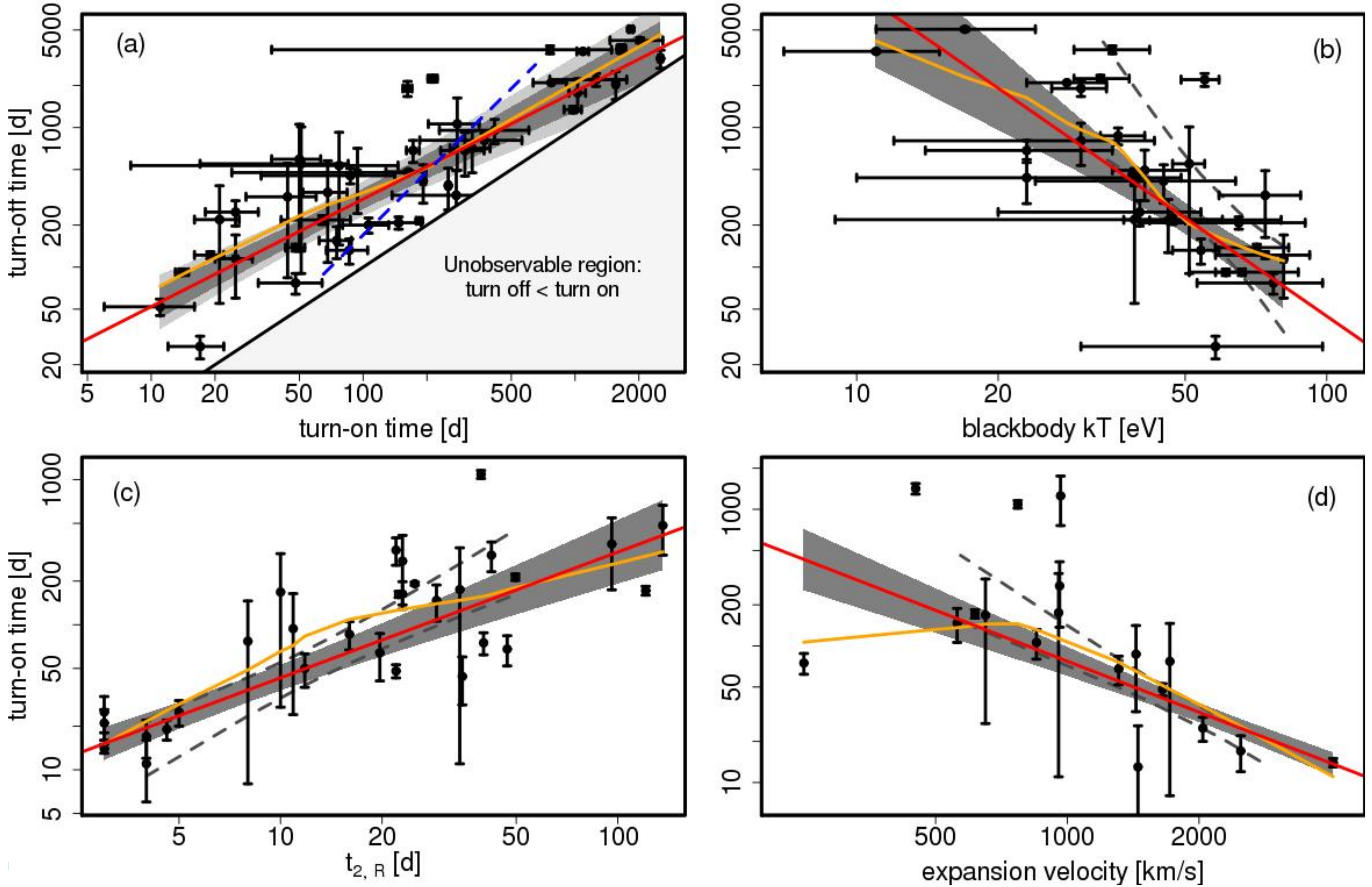
M31 nova: T_{recurr} ~ 1 yr



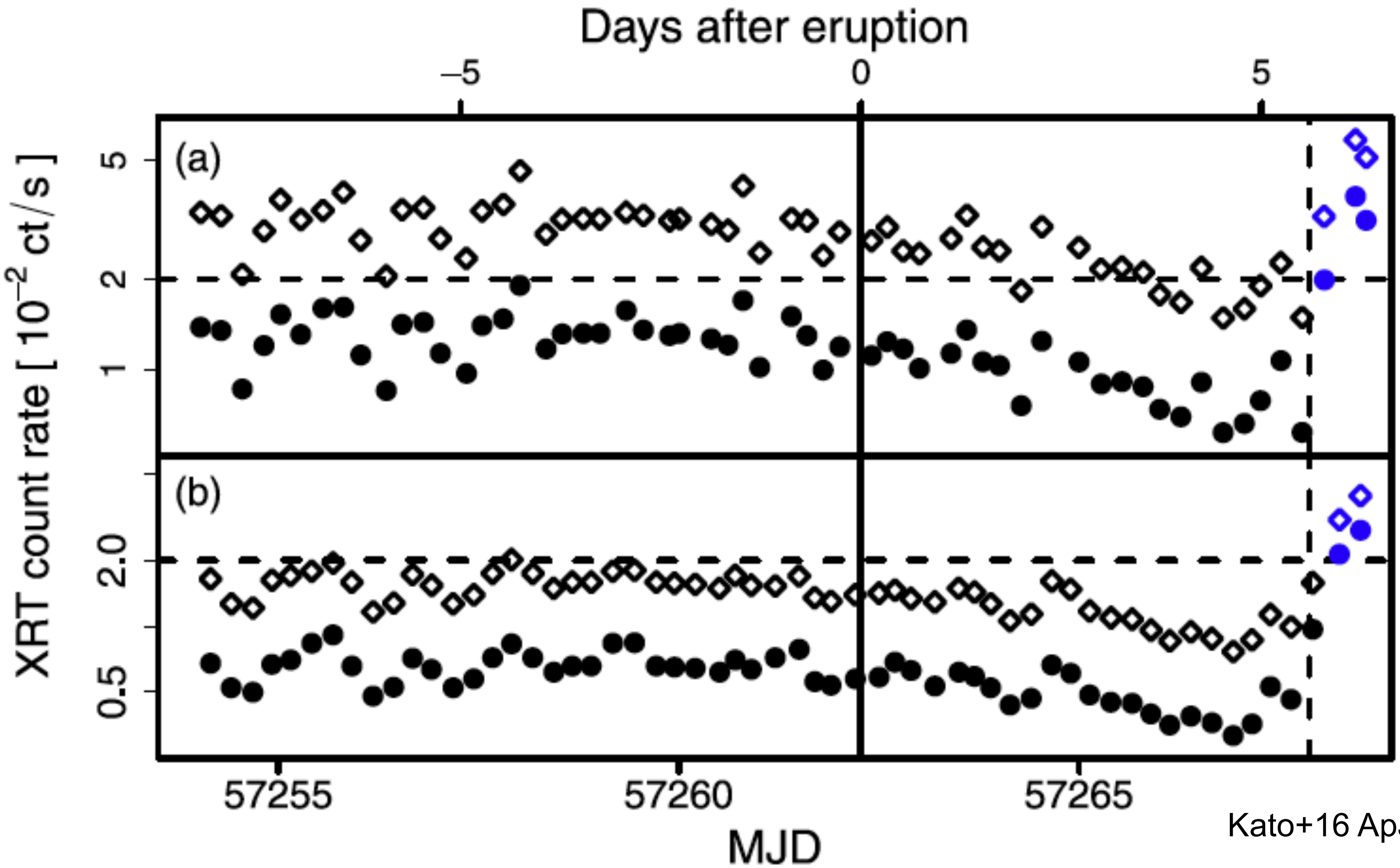
Kato+ ApJ 793, 136 (2014)

- M_{WD} > 1.3 M & acc rate > 1.5x10^-7 Msun/yr
- Min poss T_{recurr} ~ 2 months (for non-rot 1.38 M_{WD})

M31 novae: Henze+14



M31N 2008-12a



Kato+16 ApJ 830 40

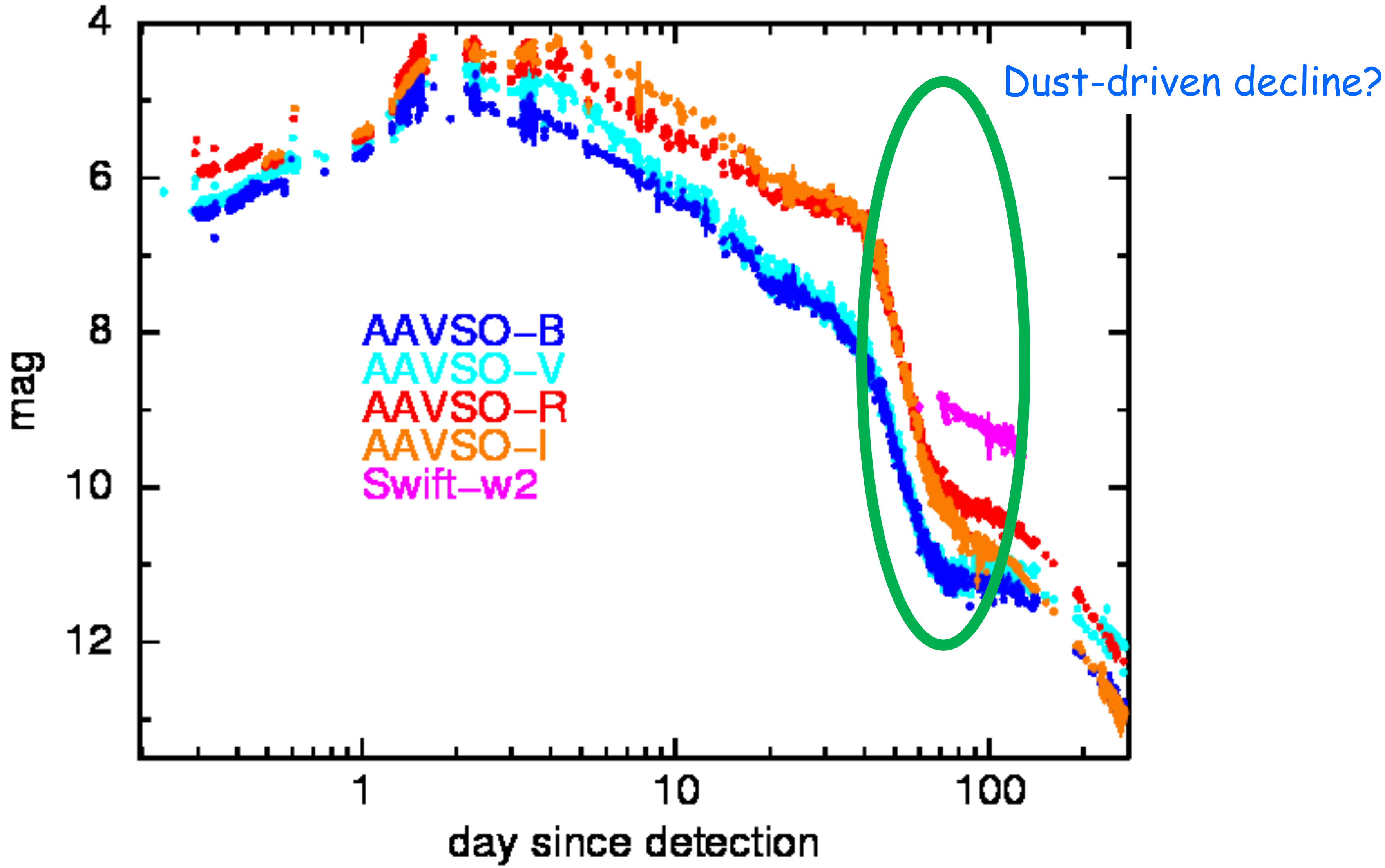
- 4 observations a day to search for shock breakout
- only upper limits found

Dust in V339 Del

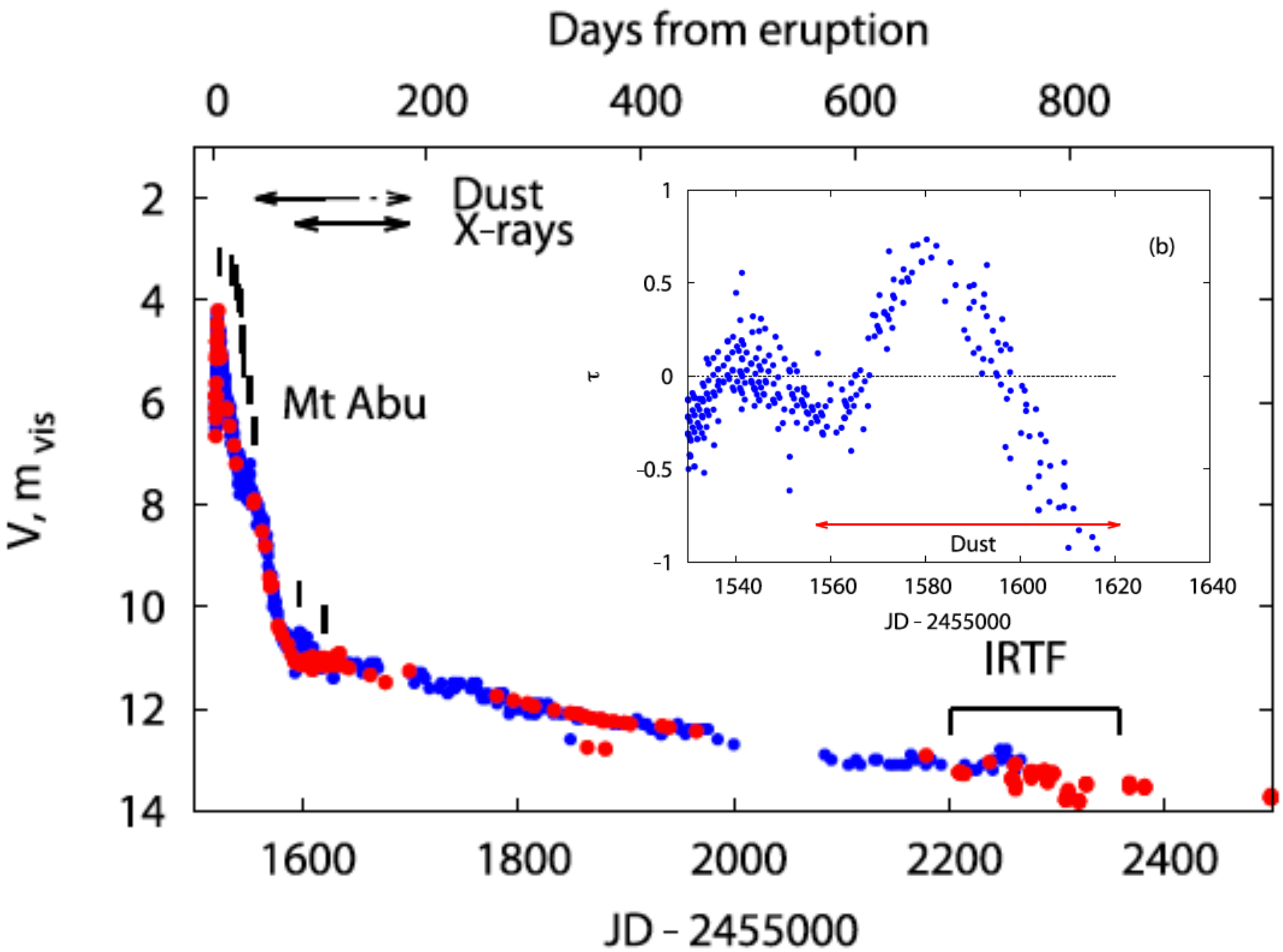
- N Del 2013 (V339 Del) was one of the brightest novae in the last century (Vmax ~ 4.3)
- LAT VHE gamma-ray source
- MS donor star, bipolar ejecta
- Filling factor ~0.1 derived from ratio of H beta luminosity to electron density from [O III]
- Ejecta mass 2-3e-5 Msun
- Taranova+14 discovered IR dust peak at ~ 1 month
- Relationship between dust & X-rays not well established

Shore+16 A&A 590 A123

Dust in V339 Del

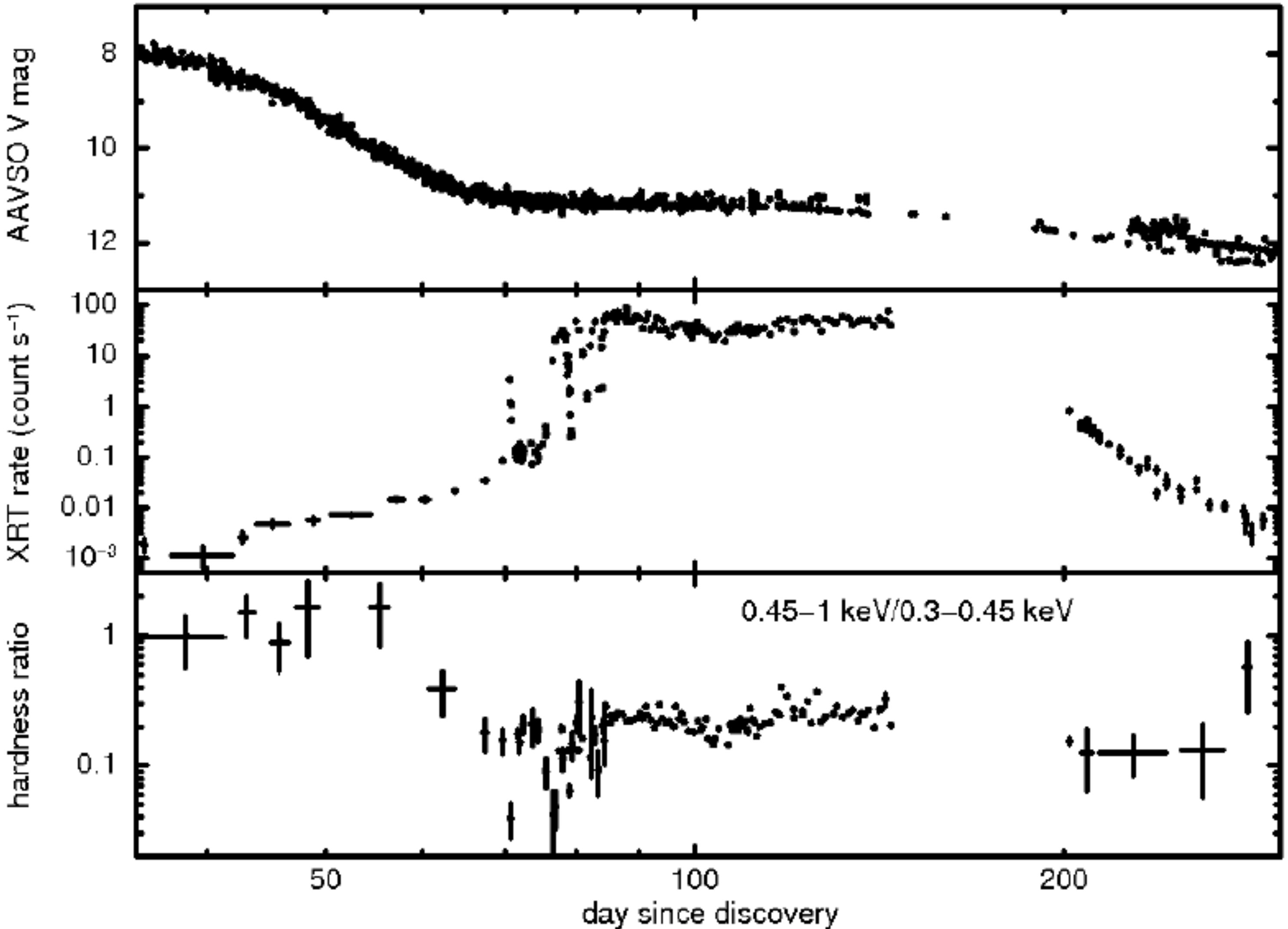


Dust in V339 Del



Evans+17 MNRAS 466 4221

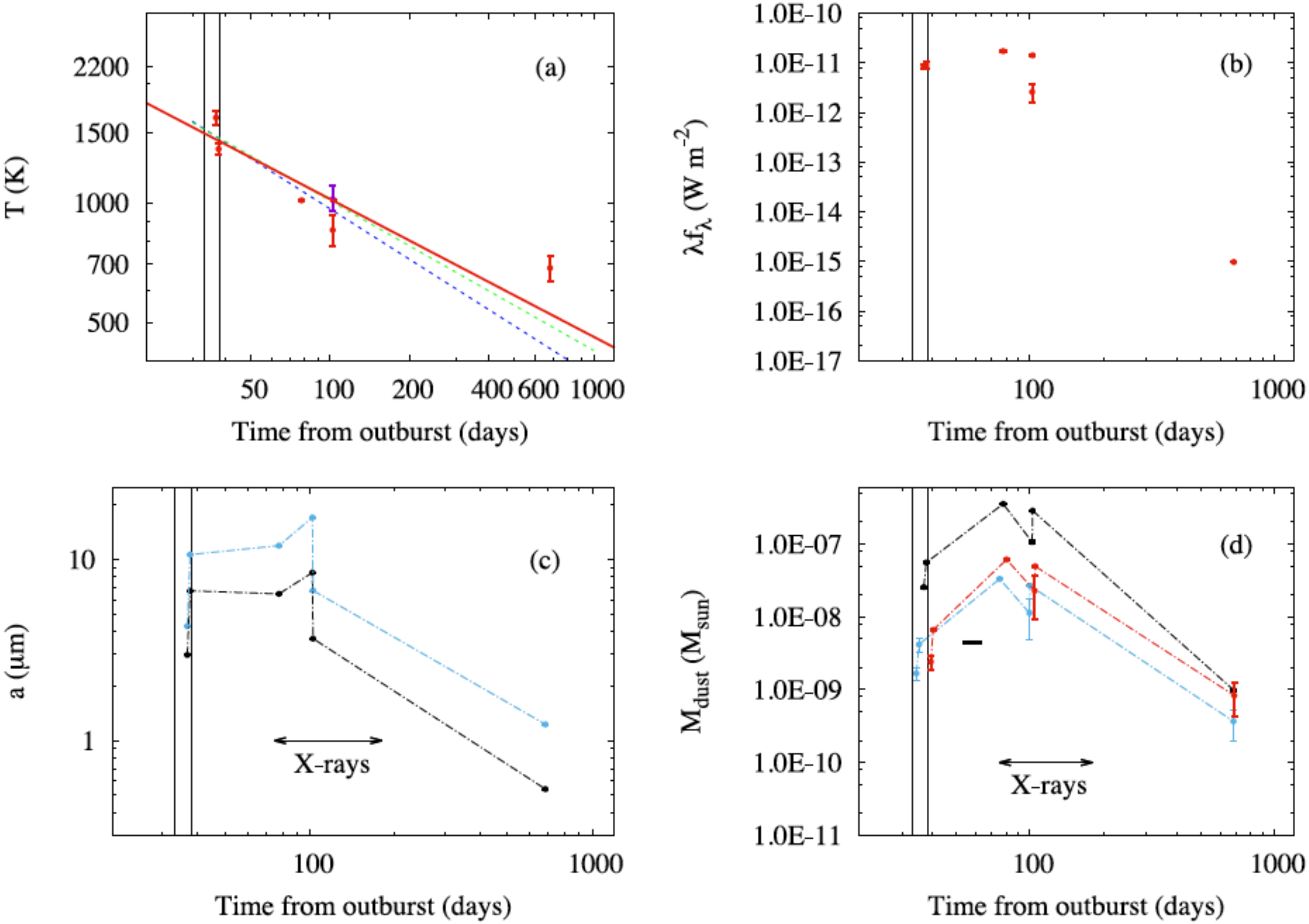
Dust in V339 Del



Shore+16 A&A 590 A123

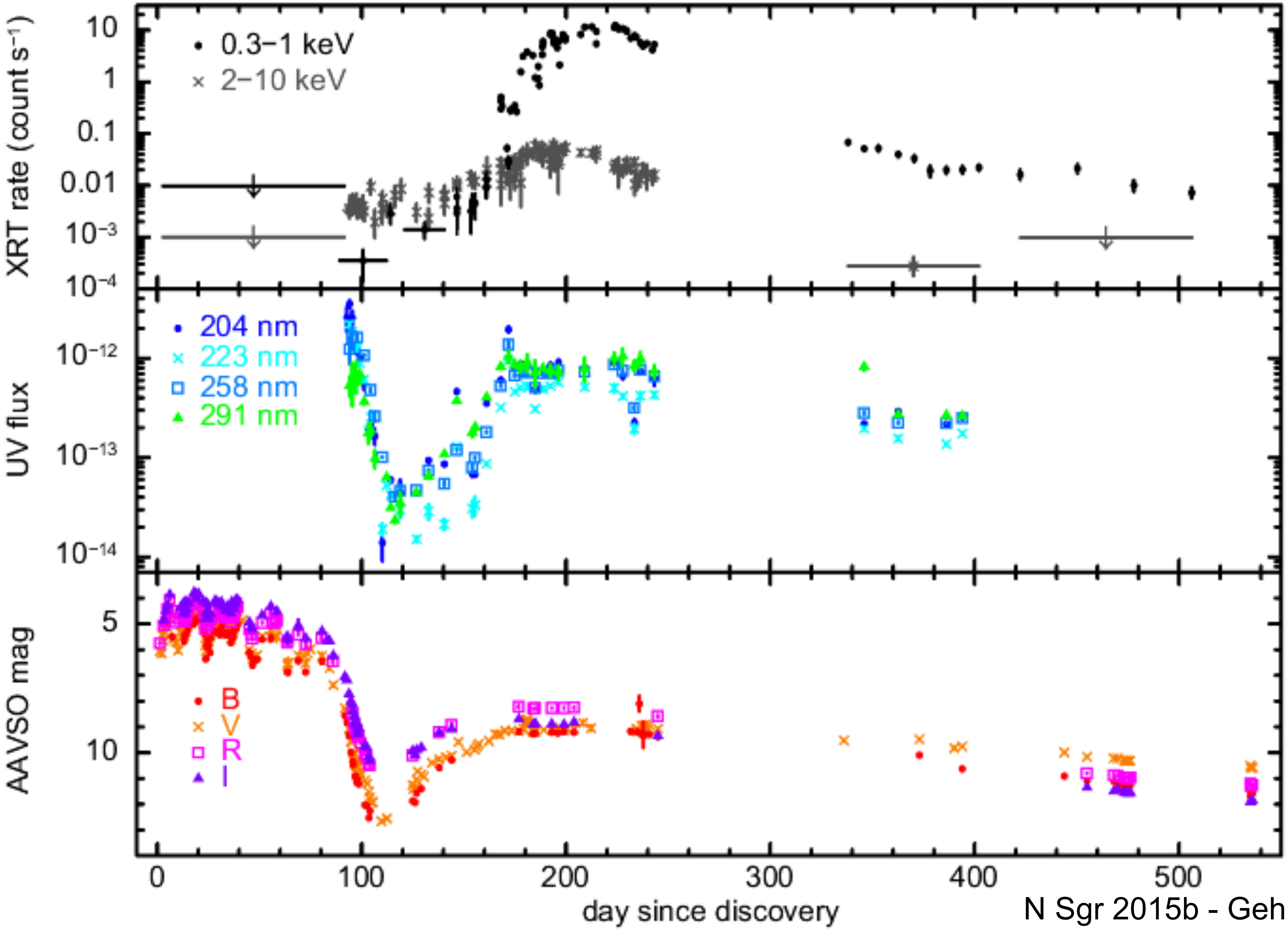
- No obvious sign of dust effect in X-ray light curve

Dust in V339 Del



Evans+17 MNRAS 466 4221

Dust in V5668 Sgr



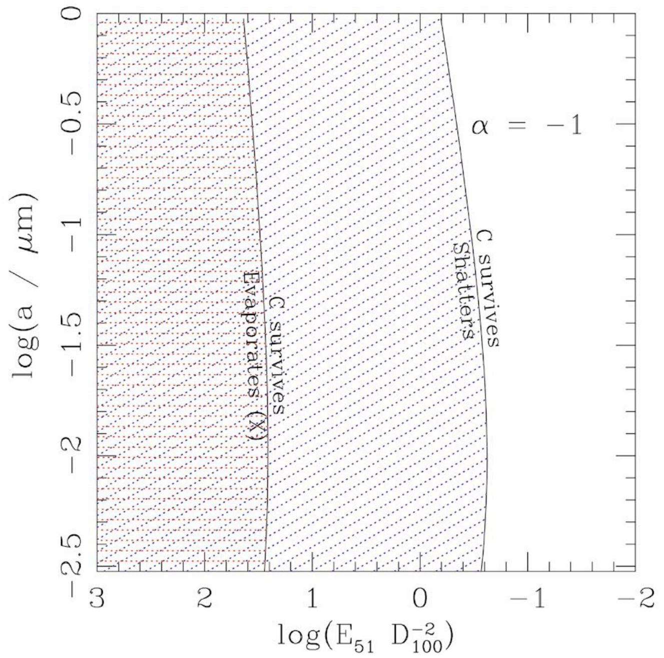
N Sgr 2015b - Gehrz+ in prep

Dust & X-rays

- CO novae may produce C dust
- C K absorption edge is at 0.27 keV
  - probably too low to be of interest
- Fructer+01 argues that dust grains smaller than ~1 micron are optically thin to X-rays
  - X-ray absorption will be unaffected by whether C is in dust or gas at these grain sizes

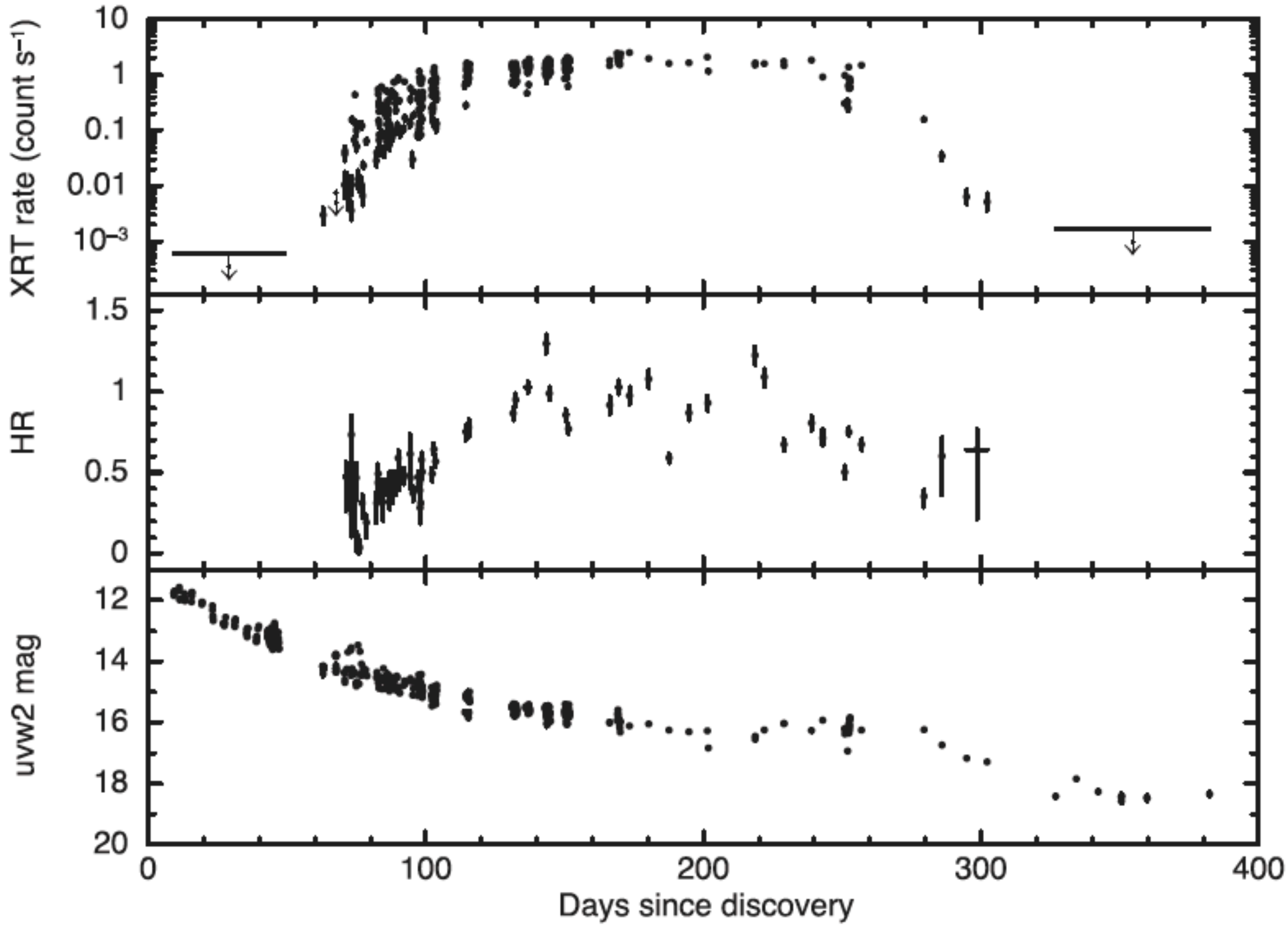
Dust & X-rays

- The observed soft X-ray rise around the times of the end of the dust events is not caused by dust destruction
- The dust destruction may be caused by the X-ray rise
- X-rays destroy dust by
  - thermal evaporation
  - electron ejection causing strong Coulomb repulsion
- V5668 Sgr has $\log(E/D^2) > 5$
- Its X-ray energy density easily destroys the grains



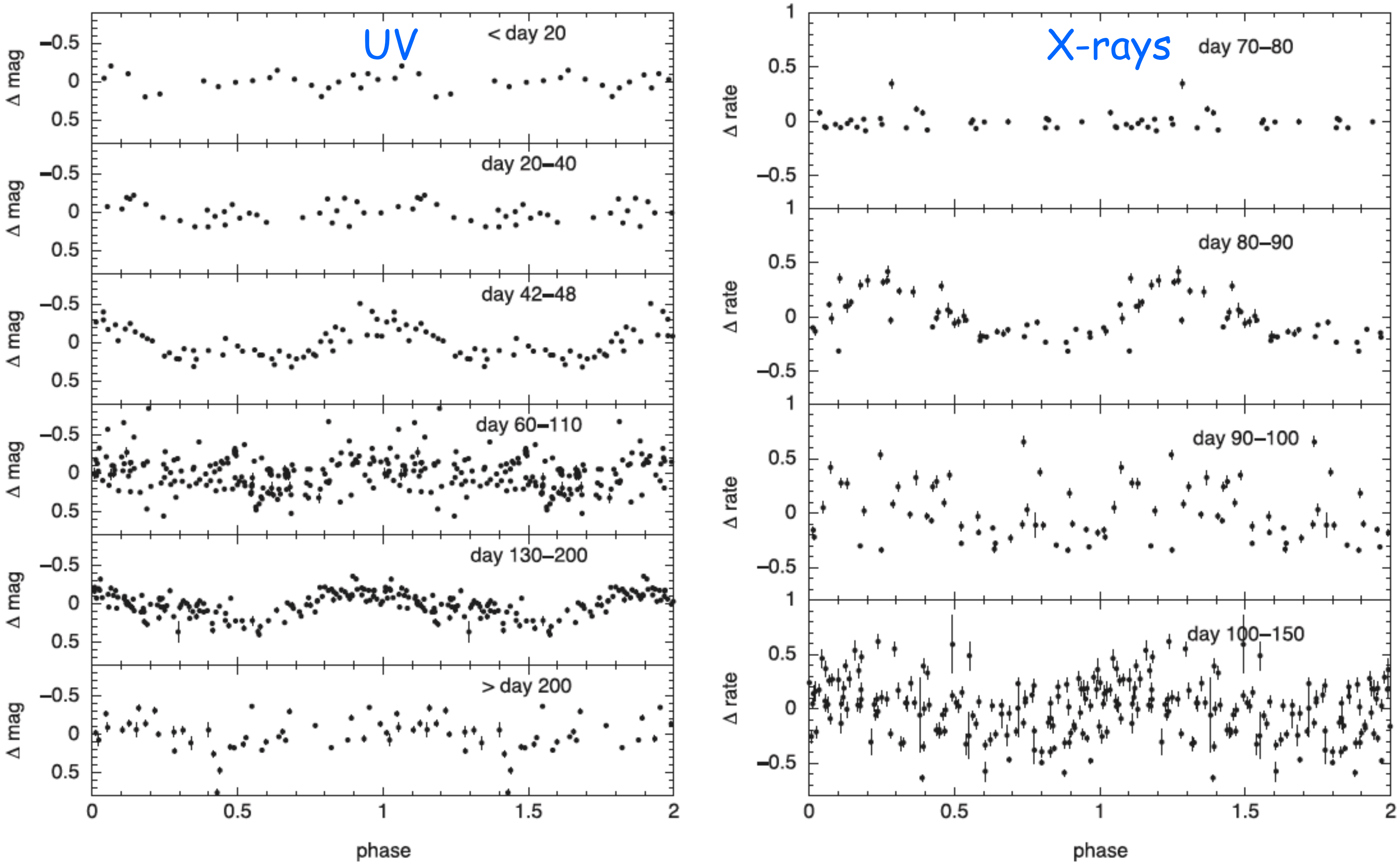
Fruchter+01

LMC 2009a



Bode+16 ApJ 818 145

LMC 2009a: 1.19d period

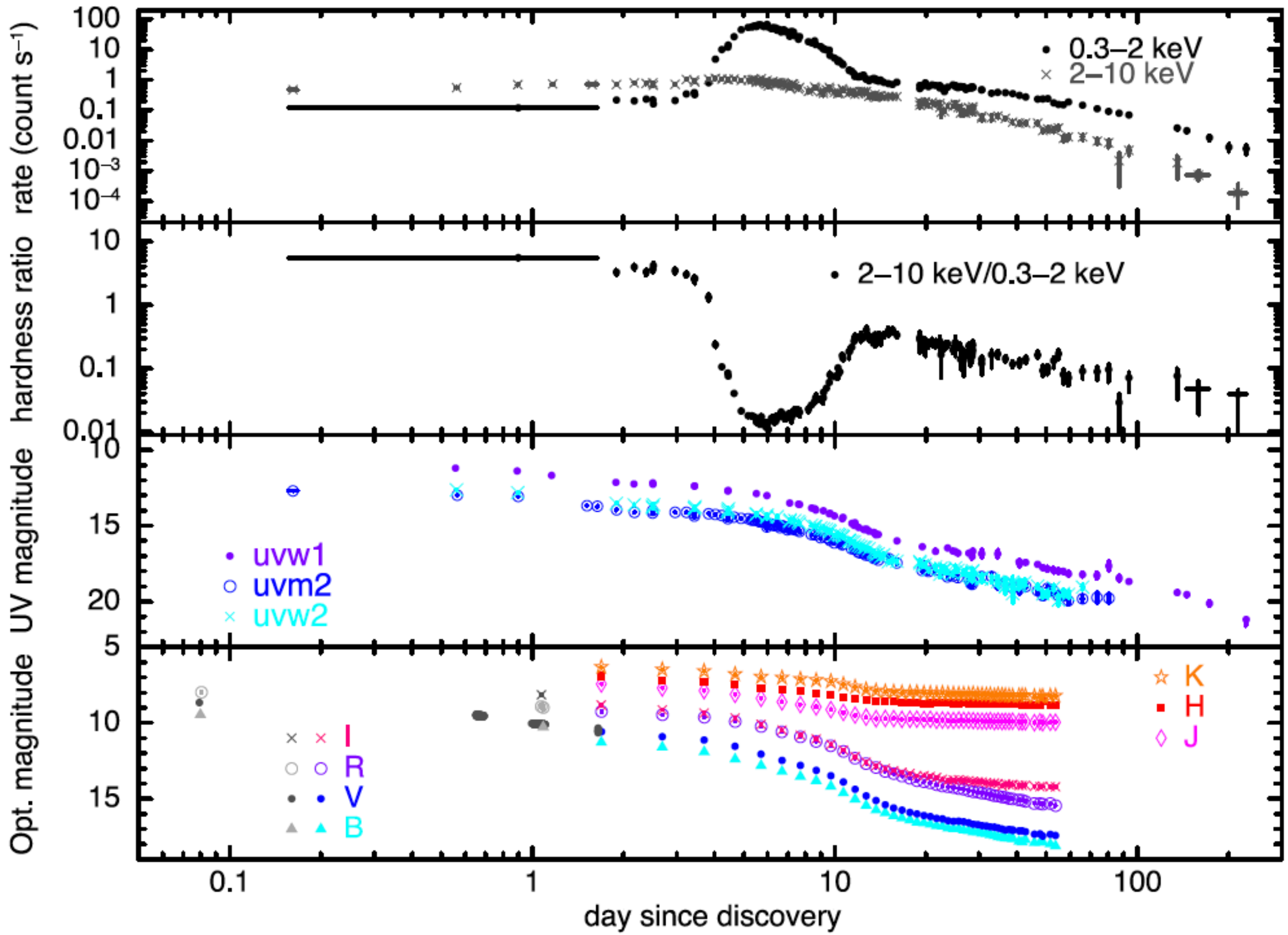


Bode+16 ApJ 818 145

LMC 2009a modulation

- 1.19 day period is too long to be WD spin
- Smooth modulation over the whole orbit implies an emitting structure size comparable to the orbit
    - scattering cloud around WD?
- UV or X-rays modulated, but not both simultaneously
- X-ray modulation lags UV by 0.28P
    - UV modulation source must be distinct from X-ray (also implied by overall light curve)
    - azimuthally structured raised disk rim?

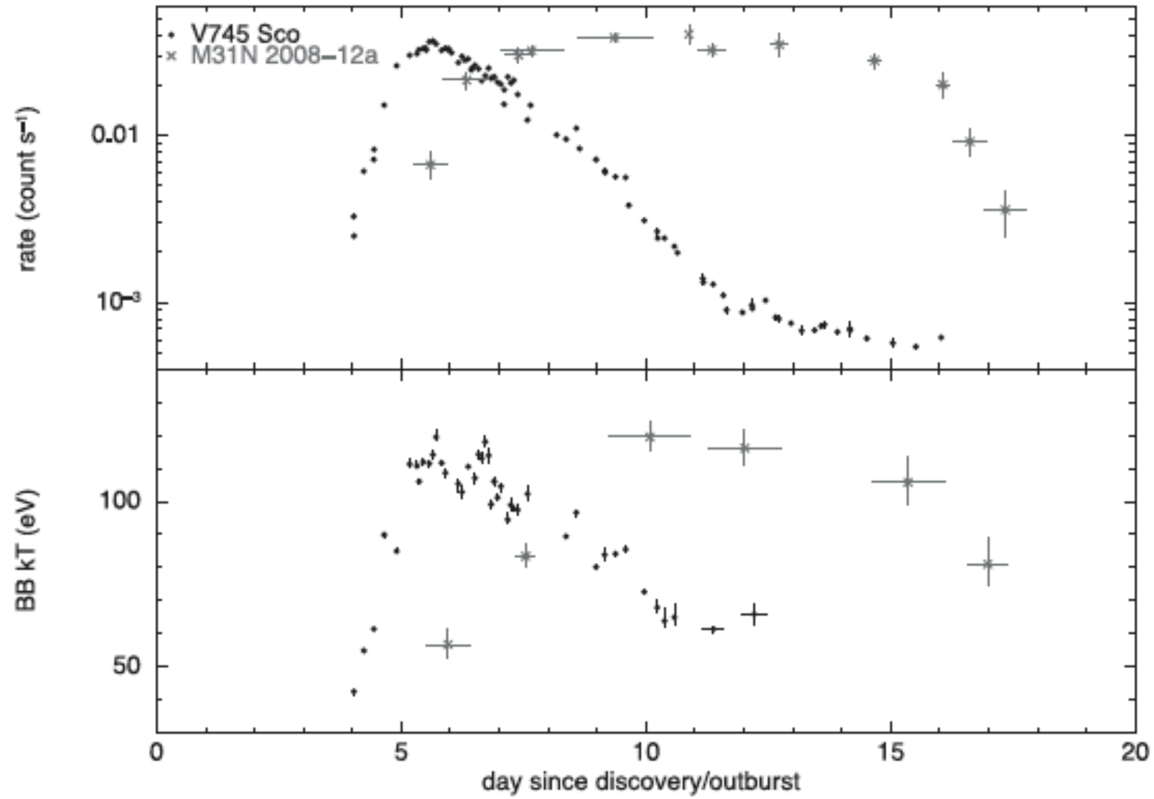
V745 Sco: Swift



- 1^{st} obsn at 3.7 hrs after discovery announcement

Page+15 MNRAS 454 3108

V745 Sco: Swift



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- V745 Sco kT 25 eV below that of M31 N (curve shifted up): WD mass lower
- Much shorter X-ray timescales than M31N 2008-12a: WD mass higher
- WD mass is not the only important parameter
- e.g. burnt & ejected mass with T_rec gives accretion rate ~5% that of M31N

V745 Sco: Swift+NuSTAR

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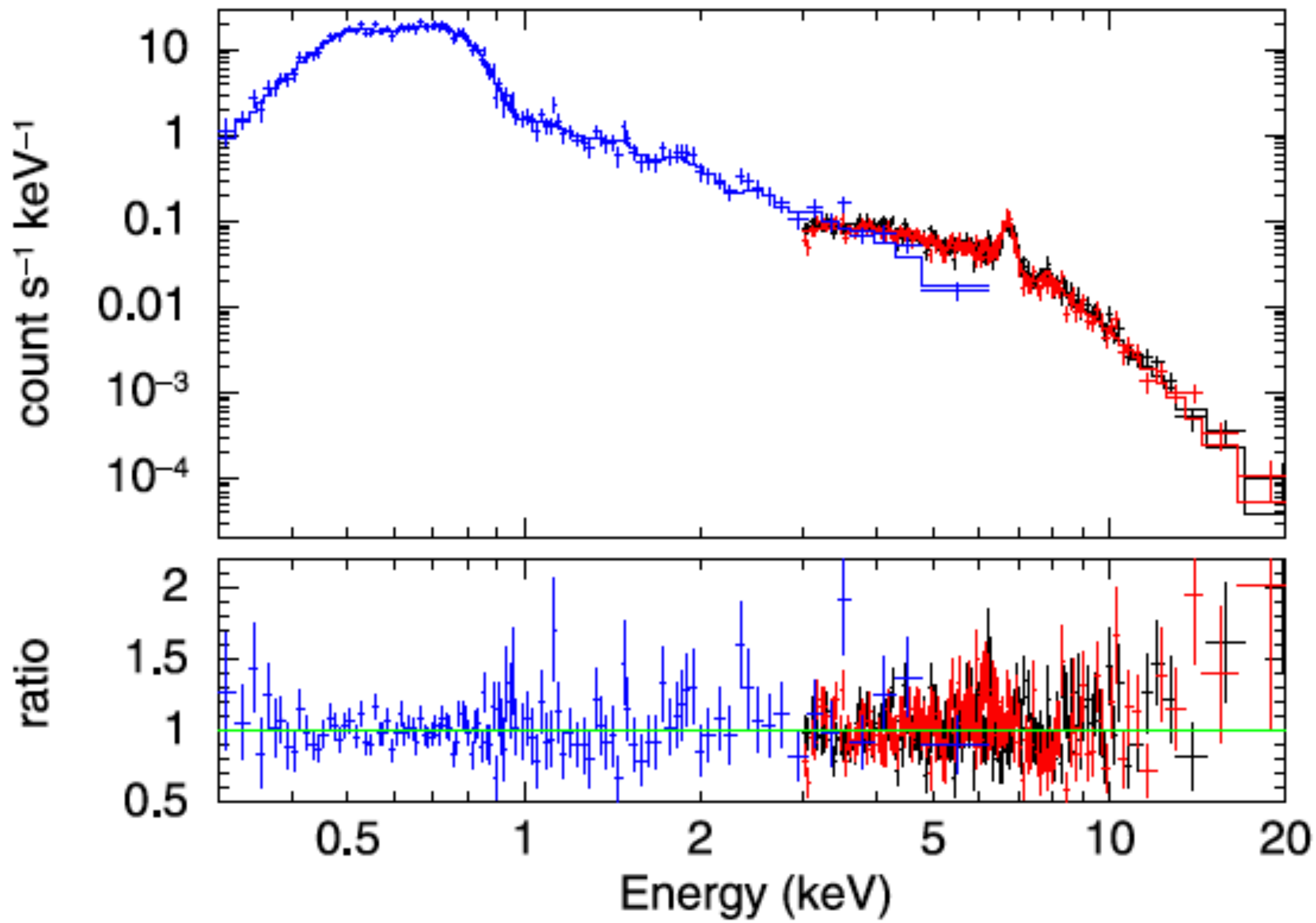
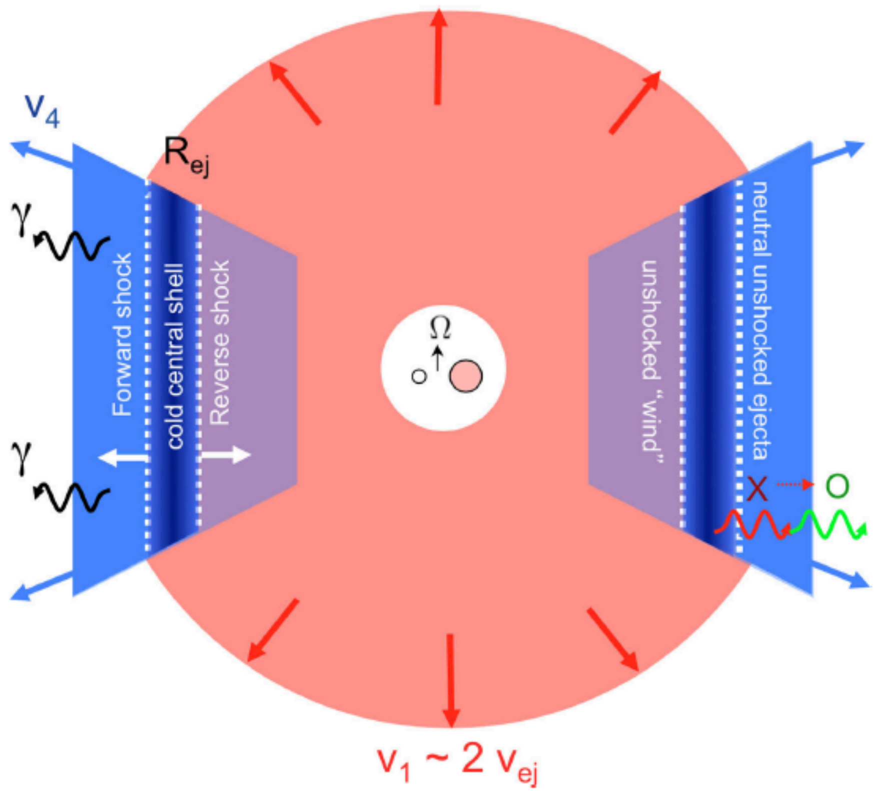


Orio+15 MNRAS 448 L35

- 2-3 sigma Fermi LAT detection on day 1-2
- Swift & NuSTAR observations on day 10

V745 Sco: Swift+NuSTAR

- Chomiuk+14 proposed that hard X-ray emission arises from shocks between fast polar and slower equatorial ejecta
- Metzger+15 proposed that significant optical emission is powered by absorbed X-ray shocks
- M+15 propose a search for these absorbed shocks with NuSTAR at the time of VHE gamma-ray emission and optical peak
- The V745 Sco observation was unfortunately too late (due to the exceptional speed of this nova)
- This test has yet to be performed AFAIK



Metzger+15 MNRAS 450 2739

Summary

- High-energy nova observations have revealed surprises
- Not really understood:
  - Large amplitude variability of early super-soft phase ('clumps'?)
  - 30-60 sec QPO (WD spin or oscillation?)
  - GeV emission origin
- Perhaps understood:
  - High mass WDs lead to fast novae
  - Irradiation can play a large role in optical light curve formation
  - Accretion disk can be present in residual nuclear burning phase
  - X-rays from inner regions can be scattered into line of sight
  - Completeness of fast novae may be very poor
- Swift has contributed answers, questions & motivation for observations elsewhere; and can continue to do so